

Continuum observations of M 51 and M 83 at 1.1 mm with AzTEC

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ABSTRACT

We observed the spiral galaxies M 51 and M 83 at 20 arcsec spatial resolution with the bolometer array Aztronomical Thermal Emission Camera (AzTEC) on the JCMT in the 1.1 mm continuum, recovering the extended emission out to galactocentric radii of more than 12 kpc in both galaxies. The 1.1 mm-continuum fluxes are 5.6 ± 0.7 and 9.9 ± 1.4 Jy, with associated gas masses estimated at $9.4 \times 10^9 M_{\odot}$ and $7.2 \times 10^9 M_{\odot}$ for M 51 and M 83, respectively. In the interarm regions of both galaxies, the $N(\text{H}_2)/\text{I}(\text{CO})$ (or X-factor) ratios exceed those in the arms by factors of ~ 1.5 –2. In the inner discs of both galaxies, the X-factor is about $1 \times 10^{20} \text{ cm}^{-2} (\text{K km s}^{-1})^{-1}$. In the outer parts, the CO-dark molecular gas becomes more important. While the spiral density wave in M 51 appears to influence the interstellar medium and stars in a similar way, the bar potential in M 83 influences the interstellar medium and the stars differently. We confirm the result of Foyle et al. that the arms merely heighten the star formation rate (SFR) and the gas surface density in the same proportion. Our maps reveal a threshold gas surface density for an SFR increase by two or more orders of magnitude. In both galaxy centres, the molecular gas depletion time is about 1 Gyr climbing to 10–20 Gyr at radii of 6–8 kpc. This is consistent with an inside-out depletion of the molecular gas in the discs of spiral galaxies.

Key words: galaxies: individual: M 51, M 83 – galaxies: ISM – galaxies: spiral.

1 INTRODUCTION

Central to a complete picture of galaxy evolution is the distribution of the interstellar matter (ISM) within each galaxy and how that ISM forms stars. Given that the ISM mass on galactic scales is dominated by molecular and atomic gas, observing the tracers of these gas components is necessary for measuring the ISM distribution within the discs of spiral galaxies. Accordingly, observations of the H I 21-cm line and the CO $J = 1 \rightarrow 0$ 2.6-mm line are often used as tracers of the atomic and molecular gas, respectively (see, e.g. Nishiyama, Nakai & Kuno 2001; Regan et al. 2001; Regan 2006; Walter et al. 2008). While conversion of the velocity-integrated brightness temperature of the H I line, or $\text{I}(\text{H I})$, to atomic gas column density, $N(\text{H I})$, is usually straightforward (though not always, e.g. Planck Collaboration XIX 2011b), the conversion of $\text{I}(\text{CO})$ to molecular gas column density $N(\text{H}_2)$ is not quite so certain (Rickard & Blitz 1985;

Israel 1988; Maloney & Black 1988; Wall et al. 1993; Sodroski et al. 1995; Dahmen et al. 1997, 1998; Regan 2000; Paglione et al. 2001; Wall 2007; Narayanan et al. 2011; Narayanan et al. 2012; Shetty et al. 2011a,b; Papadopolis et al. 2012), especially given that the CO $J = 1 \rightarrow 0$ line is known to be optically thick (e.g. see Evans 1980; Kutner 1984; Evans 1999). Recent observations (Planck Collaboration XIX 2011b) suggest that the $N(\text{H}_2)/\text{I}(\text{CO})$ conversion factor, or X-factor, X_{F} , is roughly constant within the disc of our Galaxy, with $X_{\text{F}} = (2.5 \pm 0.1) \times 10^{20} \text{ H}_2 \text{ cm}^{-2} (\text{K km s}^{-1})^{-1}$, to be abbreviated as $X_{\text{F}} = 2.5 \pm 0.1 X_{20}$; this or a similar value of X_{F} is often called the ‘standard’ value. This uniform X-factor value for our Galaxy’s disc now applies to the discs of external galaxies, where $X_{\text{F}} \sim 2 X_{20}$ is inferred and, on average, is radially non-varying from the inner disc to a galactocentric radius of $\sim 1 R_{25}$ (Sandstrom et al. 2013). The evidence for CO-dark gas, both theoretically and observationally, is a further complication (see, for example Roman-Duval et al. 2010; Planck Collaboration XIX 2011b; Clark et al. 2012; Sain-tonge et al. 2012; Langer et al. 2014; Smith et al. 2014). It is thus advantageous to employ tracers other than CO $J = 1 \rightarrow 0$ as

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independent checks on ISM surface density variations to test recent physical models of X_F (e.g. Narayanan et al. 2011, 2012; Shetty et al. 2011a,b; Papadopolis et al. 2012).

Observed gas column and surface densities can then provide insights into large-scale star formation in galaxies. The Schmidt–Kennicutt (S-K) law, for example, states that star formation rate (SFR) surface density, Σ_{SFR} , is related to the gas surface density, Σ_{gas} , by $\Sigma_{\text{SFR}} \propto \Sigma_{\text{gas}}^\alpha$ with $\alpha \sim 1.0$ to ~ 5.0 (e.g. Schmidt 1959; Kennicutt 1989, 1998; Bigiel et al. 2008; Heiderman et al. 2010; Leroy et al. 2013). The index $\alpha = 1.0$ is appealing because the gas depletion time ($\propto \Sigma_{\text{gas}}/\Sigma_{\text{SFR}}$) is constant throughout the spiral discs. While there is some evidence that $\alpha = 1$ on size scales of $\gtrsim 1$ kpc (e.g. Bigiel et al. 2008; Calzetti, Liu & Koda 2012; Leroy et al. 2013), there is also evidence of non-linear and even non-universal slopes on such size scales (see, e.g. Shetty, Kelly & Bigiel 2013; Pan, Kuno & Hirota 2014; Santini et al. 2014), especially on the scales of individual giant molecular clouds (GMC; Lada et al. 2013; Lombardi et al. 2014). There is also strong evidence for an inside-out formation of galaxies (e.g. González Delgado et al. 2014), which is at odds with a constant gas depletion time and, therefore, with having $\alpha = 1$. Comparison between those results on the large (i.e. galactic) scales with those on GMC scales is problematic. Lada et al. (2013) suggest that the Schmidt law on the scales of GMCs are fundamentally different from the S-K law apparent on larger (i.e. galactic) scales; the latter are not the ‘result of an underlying physical law of star formation.’

Many of the abovementioned results used observations of the optically thick CO $J = 1 \rightarrow 0$ line and adopted a spatially constant X_F . In contrast, millimetre (mm), submillimetre (submm), and far-infrared (far-IR) continuum observations sample optically thin continuum emission from the dust grains that pervade both the atomic and molecular gas. Recently, there have been many papers of the far-IR/submm continuum emission of external galaxies from the *Planck* and *Herschel* missions (e.g. Boselli et al. 2010; Eales et al. 2010; Gordon et al. 2010; Roman-Duval et al. 2010; Planck Collaboration I 2011a; Boselli et al. 2012; Bourne et al. 2012; Foyle et al. 2012; Fritz et al. 2012; Magnelli et al. 2012; Rowlands et al. 2012; Smith et al. 2012; Foyle et al. 2013; Cortese et al. 2012, 2014; Kirkpatrick et al. 2014). These papers find, for example, that the dust and stellar masses of galaxies are correlated (Bourne et al. 2012; Cortese et al. 2012) and that spiral galaxies and dusty early-type galaxies have $\sim 10^6$ – $10^8 M_\odot$ of dust (Rowlands et al. 2012).

A major stumbling block to determining accurate surface densities from dust continuum emission is the unknown mass absorption coefficient, κ_ν , at millimetre wavelengths. Millimetre continuum emission is less temperature sensitive than that at submillimetre and far-IR wavelengths; this provides an important constraint on dust mass and sometimes the spectral emissivity index, β , can be constrained as well. Observationally, the relevant quantity determined is the dust optical depth to gas column density ratio, $\tau_{\text{vd}}/N(\text{H})$. Planck Collaboration XIX (2011b) and Planck Collaboration XXV (2011c) have found that $\tau_{\text{vd}}/N(\text{H}) = 5.2 \times 10^{-26} \text{ cm}^2$ at 857 GHz (wavelength of 350 μm) in the H I gas in the solar neighbourhood and $\tau_{\text{vd}}/N(\text{H}) = 1.1 \times 10^{-25} \text{ cm}^2$ at 250 μm (corresponding to $\tau_{\text{vd}}/N(\text{H}) = 6.0 \times 10^{-26} \text{ cm}^2$ at 857 GHz for $\beta = 1.8$) in the H I gas in the Taurus molecular complex. Given that $\beta = 1.8$ applies to the dust in our Galaxy (see, e.g. Planck Collaboration XIX 2011b) and that the dust to hydrogen gas mass is about 0.01, those observed $\tau_{\text{vd}}/N(\text{H})$ correspond to $\kappa_\nu(1.1 \text{ mm}) \simeq 0.4$ – $0.5 \text{ cm}^2 \text{ g}^{-1}$ in the dust associated with H I. The dust associated with H₂, however, has $\tau_{\text{vd}}/N(\text{H})$ double that in H I (Planck Collaboration XXV 2011c). Consequently, estimating $\tau_{\text{vd}}/N(\text{H})$ and κ_ν from comparing the ob-

Table 1. Adopted parameter values.

Parameter	M 51	M 83
Centre position (2000.0)	13 ^h 29 ^m 52 ^s .71 ^a 47° 11′ 42″.80 ^a	13 ^h 37 ^m 00 ^s .8 ^b −29° 51′ 59″ ^b
Distance (Mpc)	8.4 ^c	4.5 ^d
Position angle	170° ^e	225° ^f
Inclination	20° ^e	24° ^f

^aHagiwara et al. (2001).

^bMiller, Bregman & Wakker (2009).

^cSee Shetty et al. (2007) and references therein.

^dKarachentsev et al. (2002).

^eTully (1974).

^fTalbot, Jensen & Dufour (1979).

served dust continuum emission against the H I gas emission alone, while useful, must be viewed with caution. The various estimates of κ_ν suggest that determining the exact total mass of dust within a galaxy is uncertain by a factor of a few.

In spite of the uncertain dust mass absorption coefficient, dust continuum emission can provide estimates of X_F in our Galaxy as well as in external galaxies (see, e.g. Israel 1997a,b; Reach, Wall & Odegard 1998; Eales et al. 2010; Roman-Duval et al. 2010; Planck Collaboration XIX 2011b; Sandstrom et al. 2013). Such observations have shown that while X_F can be more or less spatially constant in some cases, like in the disc of our Galaxy and other external galaxies (see Planck Collaboration XIX 2011b; Sandstrom et al. 2013), there can be regions of ‘dark’ gas, H₂ with no CO emission, both in our Galaxy and other galaxies (e.g. see Roman-Duval et al. 2010; Planck Collaboration XIX 2011b; Clark et al. 2012; Saintonge et al. 2012; Baes et al. 2014; Langer et al. 2014; Pineda, Langer & Goldsmith 2014; Smith et al. 2014). Therefore, observations of dust continuum emission provide a vital check on results inferred from CO $J = 1 \rightarrow 0$ observations.

1.1 The current work

Even with the many recent advances mentioned above, there are many questions left unanswered. For example, do the inferred X-factor values (i.e. Planck Collaboration XIX 2011b; Sandstrom et al. 2013) apply to the outer discs of all galaxies? Also, are there systematic differences of the X-factor between arm and interarm regions? Can previous methods of observationally inferring the dust mass absorption coefficient at millimetre wavelengths be refined? How do the answers to those questions influence the specific form of the observed S-K law in a given galaxy?

To address these questions and to better understand the gas and dust in spiral galaxies and their relationship to star formation, we observed the grand-design, face-on spiral galaxies M 51 and M 83 with the bolometer-array camera, AzTEC (Aztronomical Thermal Emission Camera), mounted on the 15-m James Clerk Maxwell Telescope (JCMT) in Hawaii at a wavelength of 1.1 mm. Both of these galaxies are nearby with distances of less than 10 Mpc (see Table 1 for details) and, hence, the spiral arms in both galaxies are resolved across the arms in the JCMT/AzTEC observations, which have a spatial resolution of 20 arcsec. Both galaxies have been studied extensively at numerous wavelengths (e.g. Rots et al. 1990; Helfer et al. 2003; Kennicutt et al. 2003; Dale et al. 2009; Tilanus & Allen 1993; Crosthwaite et al. 2002; Blasco-Herrera et al. 2010; Hughes 2013; Meidt et al. 2013; Schinnerer et al. 2013; Colombo et al. 2014).

Recent *Herschel* observations of M 51 and M 83 at 70, 160, 250, 350, and 500 μm have provided maps of the dust temperature,

surface density, and even spectral emissivity index, β (Bendo et al. 2012; Cooper et al. 2012; Foyle et al. 2012, 2013), but with spatial coverage that is slightly more limited than those of the AzTEC 1.1 mm continuum maps presented here; the AzTEC images cover a few more kiloparsecs at the adopted distances given in Table 1. As a result, these AzTEC 1.1 mm continuum images extend both the spatial and the wavelength coverage of the dust emission in both M51 and M83.

This greater spatial coverage placed restrictions on the other wavelengths available for comparison with the AzTEC 1.1 mm data; there were no 250, 350, and 500 μm data towards the outer discs of M51 and M83. There are, however, *Spitzer* 160 μm data covering both of these galaxies. This wavelength is the longest of the *Spitzer* data and, along with the AzTEC 1.1 mm data, is the one most likely associated with the dust component(s) that dominate the mass of the dust.

In addition, the surface densities in the current work were estimated from the observational data. Specifically, $\tau_{\text{vd}}/N(\text{H})$ at $\lambda = 1 \text{ mm}$ was inferred by comparison with the H I column densities and removing upper outliers, because these upper outliers were assumed to represent positions with CO-dark gas. This approach has the advantage that final gas masses inferred were not dependent on dust models.

2 OBSERVATIONS

Both M51 and M83 were observed during the nights of 2005 December 6–12 and 2006 January 12–20 on the JCMT with AzTEC, which is a bolometer array for observing continuum emission at 1.1 mm (Wilson et al. 2008). AzTEC has 144 silicon nitride micromesh detectors arranged hexagonally in six ‘hextants’. The AzTEC foot-print on the sky while mounted on the JCMT was ~ 280 arcsec wide, where each detector had an 17–18 arcsec almost circular beam. During the observations, 107 of the 144 AzTEC detectors were fully functional, the faulty detectors found mostly in Hextants 1 and 2 (see fig. 11 of Wilson et al. 2008). The images of M51 and M83 were created by raster scanning, so the missing detectors did not affect the coverage of the final images, only their sensitivities.

The fields for both galaxies were originally chosen to be centred on each and were 14 arcmin \times 14 arcmin in size. This field size is large enough to accommodate each galaxy with a ~ 10 arcmin diameter, and one-half AzTEC foot-print on each side. The actual observations, however, covered nearly 25 arcmin \times 25 arcmin for each galaxy. This allows the AzTEC routines to properly remove the large-angular scale atmospheric emission from the M51 and M83 images. Both M51 and M83 were raster scanned for many cycles; the total integration time was about 12 h for M51 and about 14.5 h for M83. Given the long-term stability of the detectors, chopping was not necessary. During the observations, the zenith optical depth at 225 GHz was between about 0.05 and 0.15. The source elevations during the observations were between about 30° and 60° for M51 and between about 30° and 45° for M83. These imply that the maximum line-of-sight optical depth at 225 GHz was typically much less than 0.3.

Interleaved between groups of raster scans of the program galaxies were observations of additional sources to focus the AzTEC camera, to check the pointing, and to calibrate the data. Focusing was achieved through repeated jiggle-maps of a few chosen point-like sources to minimize the beam’s angular size and maximize its peak. Focusing was done each night usually on the planet Uranus, but also on the late-type stars CRL618 and IRC+10216.

Pointing was checked and corrected by small jiggle-maps of QSOs 1308+326 and 1334–127. The pointing maps were performed before and after many raster scans of the program galaxies, or about every 1–1.5 h. The rms pointing uncertainty was 2 arcsec or better. Beam maps of Uranus were made to determine the flux conversion factor which converts the detector output voltages to mJy beam^{−1} (see Wilson et al. 2008, for details).

Spectral lines do not contribute appreciably to the AzTEC 1.1-mm bandpass. The strongest line in this bandpass is CO $J = 2 \rightarrow 1$. This is at the edge of the bandpass for the 1.1-mm filter (see fig. 4 of Wilson et al. 2008) and line peak would be attenuated by a factor of ~ 100 due to the low response of the filter at this frequency. In addition, the AzTEC bandwidth at 1.1 mm is about 70 GHz, resulting in a spectral dilution of the line of a factor of $\sim 10^3$. The two effects together dilute the line strength by a factor of $\sim 10^5$. Publicly available CO $J = 1 \rightarrow 0$ maps (e.g. Crosthwaite et al. 2002; Helfer et al. 2003) along with adopting a reasonable ratio for the $J = 2 \rightarrow 1$ to $J = 1 \rightarrow 0$ lines (i.e. 0.7) and applying the peak attenuation, the spectral dilution, and the observed 1.1-mm fluxes of M51 and M83 (see Section 4.1) yield a relative contribution of $5\text{--}6 \times 10^{-4}$ by the CO $J = 2 \rightarrow 1$ line to the total observed 1.1-mm flux. Including the effects of other spectral lines, even those with peak attenuations closer to unity, are unlikely to add a total flux contribution of more than a few per cent to the AzTEC 1.1-mm band.

3 DATA REDUCTION

The emission observed by the telescope in the 1.1 mm continuum is dominated by that of the atmosphere. In fact, such atmospheric emission is roughly a factor of $1\text{--}3 \times 10^5$ stronger than that from astronomical sources. Hence, considerable data processing is necessary for extracting the faint astronomical signal from the time series data that will be converted into an astronomical image. This processing assumes that the spatial extent of the atmospheric emission is greater than that of the astronomical sources; this makes it possible, though still difficult, to reconstruct extended structures in the image. Recovering the extended astronomical emission requires an iterative procedure whereby the image from the first iteration is subtracted from the time series data and these image-subtracted time series data are used to construct the next iteration image, which is again subtracted from the time series data and so on, until a suitable convergence is achieved. The algorithm described here is very similar to that for the M33 observations by Komugi et al. (2011), but with more emphasis on recovering the large-scale structure.

The data reduction starts with the full pipeline routine that removes spikes and calibrates the data, and applies principal component analysis (PCA) to filter out extended emission that largely represents the earth’s atmosphere. Included with the software pipeline output were noise maps that were created by the jack-knifing technique as described in Komugi et al. (2011). Iterations or loops of the algorithm, Flux Recovery Using Iterative Technique or FRUIT, are then applied to the preliminary map in order to recover the missing extended emission. The FRUITloops subtract a preliminary map from the time series data and a new map is created from those data. The different iteration maps are examined to find the latest iteration that is free of artefacts, such as large patches of negative emission. This is usually the final iteration of FRUITloops. If no iteration map is acceptable, then the parameter values used in the full pipeline and FRUITloops are changed and those routines are run again.

Estimating reasonable parameter values for a reliable reconstruction of the image required creating simulated images. The map

from the final iteration of FRUITloops was subtracted from the time-stream data and an input test map was then added into that time-stream. The test map for M 51 was the 24 μm map of M 51 from Dale et al. (2009) and for M 83, it was the 5 GHz map of M 83 from Neininger et al. (1993). These test maps were placed on the same pixel grid as the corresponding AzTEC maps and were convolved to the 18 arcsec resolution of those maps. After being added into the time-stream data (now more or less devoid of the AzTEC detected astronomical flux), the data are processed through the full pipeline and then FRUITloops. For each iteration of FRUITloops, the difference between the output map and the input map is used to determine a reduced chi-square, χ_v^2 . If the new simulation output map is acceptable, then the data processing is re-run on the real data with the new parameter values that were used in the latest simulation. The new real-data map is removed from the time-stream data and the simulations are run again. If, however, there are image artefacts in the simulated maps or the χ_v^2 is too large, then the parameter values used in the full pipeline and in FRUITloops are changed and the simulation is re-run. This process continues, going from simulations to real-data processing and back again, until the real and simulated output maps are free of artefacts and the χ_v^2 from the simulations is around 1–2.

The final AzTEC maps were converted to units of MJy sr^{-1} and smoothed to a final resolution of 20 arcsec so that the final maps were less choppy. The conversion factor to units of mJy beam^{-1} is 8.63 for the original 18 arcsec beam.

With the help of the simulations, the optimum, or near-optimum, parameter values were determined in the processing of the M 51 and M 83 data sets. The simulations showed a minimum $\chi_v^2 = 2.0$ and $\chi_v^2 = 0.45$ for M 51 and M 83, respectively. The final map of noise levels, the sigma map, of M 51 was scaled upward by $\sqrt{2}$ to account for the χ_v^2 that was greater than unity. The large-scale emission was faithfully recovered with some minor problems for both galaxies. By differencing the final output map with the initial input map of the simulation, and scaling from simulation output map to real output map, the large-scale surface brightness of the real map was checked.

The simulations tell us that the AzTEC M 51 map underestimates the true surface brightness at an average of 0.3 MJy sr^{-1} , or the rms noise level for much of the map. This has a trivial effect locally, but has a non-trivial effect on the large scale. Accordingly, the offset determined here was not added to the final M 51 map, but the simulation input and output maps were still used to check any results derived from the M 51 AzTEC map.

The case of M 83, however, was quite different. Averaged over the entire central 12.2 kpc radius, the true surface brightness is 0.05 MJy sr^{-1} or 0.25σ lower than that of the map. This correction is small enough that using the simulation to correct the results derived for M 83 was unnecessary.

The input and output maps of simulations for M 83 were also compared visually and no strong artefacts were found.

The final maps are shown in Figs 1 and 2. Comparisons with images at visible wavelengths are given in Figs 3 and 4.

4 RESULTS

4.1 Surface brightness distribution and flux

As is clearly seen in both Figs 1 and 2, both galaxies have extended low-level emission in addition to two prominent spiral arms. In the M 83 image, the structure visible within a 1 kpc radius of the centre represents M 83's bar. The total derived fluxes are $5.6 \pm 0.7 \text{ Jy}$ for

M 51 and $9.9 \pm 1.4 \text{ Jy}$ for M 83. (See Appendix A for more details.) For the adopted distances, the luminosity at 1.1 mm is $L_\nu(1.1 \text{ mm}) = (4.7 \pm 0.6) \times 10^{29} \text{ erg s}^{-1} \text{ Hz}^{-1}$ or $\nu L_\nu(1.1 \text{ mm}) = (3.1 \pm 0.4) \times 10^7 L_\odot$ for M 51. For M 83, these are $L_\nu(1.1 \text{ mm}) = (2.4 \pm 0.3) \times 10^{29} \text{ erg s}^{-1} \text{ Hz}^{-1}$ or $\nu L_\nu(1.1 \text{ mm}) = (1.7 \pm 0.2) \times 10^7 L_\odot$.

On the large scale, each galaxy possesses an exponential disc, as is seen in Figs 5 and 6, where the non-trivial deviations from the exponential fit represent the overlying spiral structure. The exponential scalelengths are given in Table 2 are given for different tracers of galactic structure, including that from the current work. For M 83, the CO $J = 1 \rightarrow 0$, near-IR, and 1.1-mm continuum have scalelengths that are comparable to within 1 or 2σ . For M 51, however, the 1.1 mm exponential disc is more than $5\text{-}\sigma$ larger than both the stellar disc (as represented by the near-IR) and the disc of CO $J = 1 \rightarrow 0$ emission. Even excluding the correction factor derived from the simulations would reduce this difference by very little. Meijerink et al. (2005) estimated a scalelength for the dust from 850 μm observations that is consistent with 1.1-mm observations.

The 1.1 mm maps of M 51 and M 83 are very similar to their corresponding CO $J = 1 \rightarrow 0$ and H I 21-cm maps, as we can see in Fig. 7. For example, the CO maps of both galaxies are similar to that of the 1.1 mm continuum out to a radius of about 6 kpc. Beyond that radius, the millimetre-continuum emission extends further than the CO $J = 1 \rightarrow 0$ emission, especially in the northern arm of M 51. That more extended emission is partly due to dust associated with atomic gas. In M 83, there is a bridge of millimetre-continuum emission extending out to 9 arcmin or about 12 kpc from the nucleus to both the southeast and to the north-west. This bridge is also seen in H I, but is shifted counterclockwise in position-angle by about 15° with respect to the millimetre-continuum bridge, where this shift is likely an artefact of the missing large-scale H I emission in the interferometer map.

4.2 Surface density distribution and mass

The 1.1 mm continuum surface-brightness maps of Figs 1 and 2 can be converted to maps of surface density, or column density, if the dust temperature is known at each position. Hence, the AzTEC 1.1 mm maps were ratioed with the *Spitzer*/MIPS 160 μm , effectively 155.9 μm , maps to yield these temperatures. The derived column densities were calibrated against the H I gas column densities at those positions where molecular appears to not dominate. This approach was modified by removing upper outliers resulting in the intermediate- κ_ν case adopted for the current paper. (See Appendix B for details.)

Figs 8–11 display the derived dust temperature maps, the column density maps, and their radial profiles for both M 51 and M 83. Both the radial T_d profile of M 51 in Fig. 9 and that of M 83 in Fig. 11 have relatively constant temperatures of $\sim 20\text{--}25 \text{ K}$ out to a radius of 3–4 kpc, a linear decline out to 15 kpc for M 51 and 9 kpc for M 83, and then a more or less constant temperature of $\sim 12 \text{ K}$ in the outer disc.

Foyle et al. (2012) had data at five wavelengths and could produce maps of both dust temperature and β . The apparently lower dust temperature in the spiral arms compared to that of the interarm dust is merely an artefact of not accounting for the spatial variation of β . In general, the dust temperatures determined from the 1.1 mm data are $\sim 5 \text{ K}$ lower than those determined from the shorter wavelengths of the *Herschel* data. This suggests that the millimetre continuum is sampling an additional component of the dust. The dust mass derived in the current work, using their dust-to-gas mass ratio and dust mass absorption coefficient (equivalent to our low- κ_ν case) is a

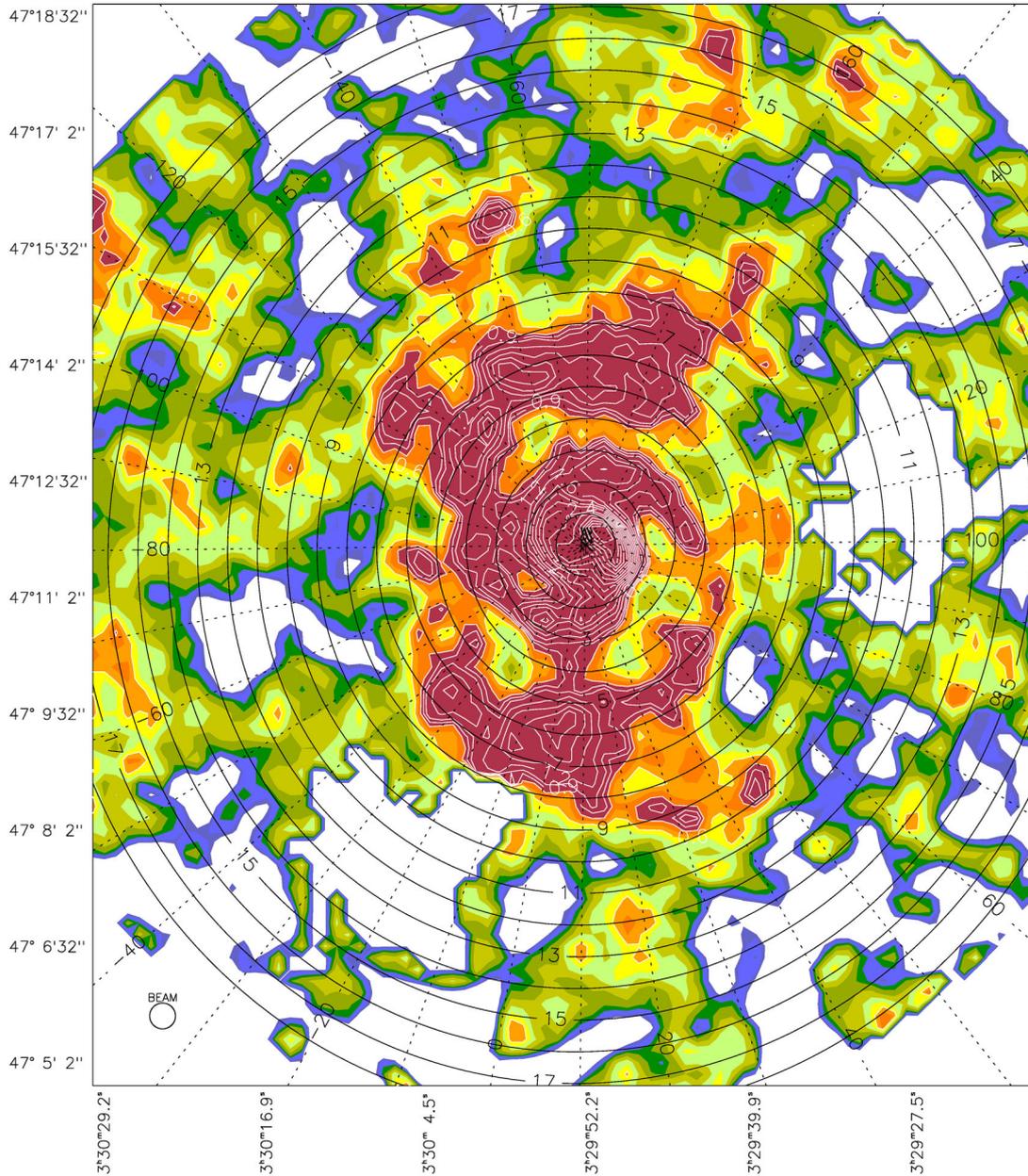


Figure 1. M51 continuum map at wavelength 1.1 mm. The coordinates are epoch 2000.0. The solid white line contour levels are 0.6, 0.8, 0.9, 1.0, 1.2, 1.4, . . . , 3.8 MJy sr⁻¹. The concentric ellipses represent distances in 1 kpc increments from the adopted centre position in the plane of M51 for its adopted distance. The dashed lines represent position angles relative to the major axis. See Table 1 for details. The effective beam is shown in the lower-left corner.

factor of 2.5 higher than their $4 \times 10^7 M_{\odot}$ over the equivalent area. This is roughly consistent with the lower T_d values that we derive. Also, based on the AzTEC/*Spitzer* data alone, the area covered by the Foyle et al. (2012) map is sampling half the dust and gas mass of M83.

A similar comparison between the *Herschel* M51 observations (Cooper et al. 2012) and those of AzTEC yield nearly identical masses (<1 per cent difference) for the Cooper et al. (2012) NGC 5194 field, after adjusting to their adopted distance, to the low- κ_v case, and to their gas-to-dust mass ratio. Repeating this for their NGC 5195 field, however, yields a disagreement of a factor of 4. This can be at least partly attributed to the differences between their observed dust temperatures and ours. They find $T_d \sim 35$ K, whereas we find $T_d \sim 22$ K. These temperatures roughly correspond to the AzTEC/*Spitzer* data yielding a dust mass that is

a factor of 4 higher. As was the case for M83, the longer wavelength AzTEC observations appear to be sampling an additional component. Interestingly, this does not seem to be the case for the NGC 5194 field, where agreement is very tight. Also, based on the AzTEC/*Spitzer* data alone, the area covered by the Cooper et al. (2012) map is sampling one-third the dust and gas mass of M51. If the constant-offset correction determined from the simulations is not applied, then this one-third is a less extreme one-half.

Despite the differences found for each of the M83 and for the NGC 5195 fields, it must be emphasized that the NGC 5194 field has the same mass for both data sets (after applying the appropriate corrections mentioned previously); this strongly suggests that the millimetre-wavelength data are not always necessary for a reasonable dust mass estimate and also that having only two continuum wavelengths will give reasonable mass estimates (i.e.

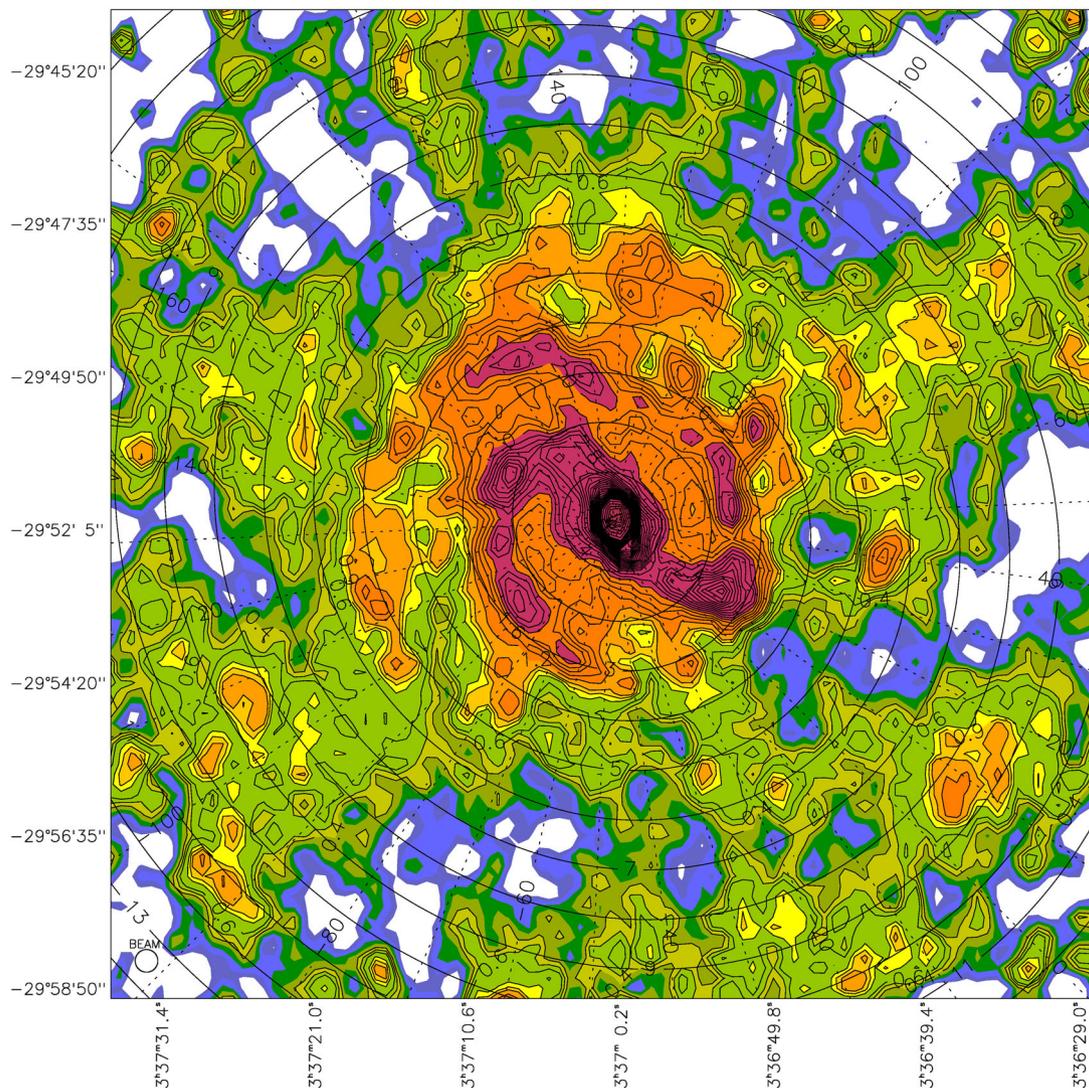


Figure 2. M 83 continuum map at wavelength 1.1 mm. The coordinates are epoch 2000.0. The solid line contour levels are 0.4, 0.5, 0.6, 0.8, 0.9, 1.0, 1.2, 1.4, . . . , 7.2, 8.0, 9.0, 10.0, 11.0, 12.0 MJy sr⁻¹. The concentric ellipses represent distances in 1 kpc increments from the adopted centre position in the plane of M 83 for its adopted distance. The dashed lines represent position angles relative to the major axis. See Table 1 for details. The effective beam is shown in the lower-left corner.

dust-inferred gas mass), especially when calibrated against gas column densities.

It should also be mentioned that the companion galaxy, NGC 5195, contributes very little to the total masses or fluxes: ~ 5 per cent. Its presence can be largely ignored.

Figs 9 and 11 demonstrate detailed agreement to within factors of about 2 between the spectral line derived and continuum derived column densities out to a galactocentric radius of about 6 kpc for both galaxies. Beyond a radius of about 8 kpc, the spectral line derived column densities are a factor of 2 or more below those derived from the continuum. This is possible evidence that the X-factor is strongly rising in the outer discs of both galaxies, although other explanations are not entirely ruled out. This will be addressed in more detail in Sections 4.3 and 5.1.

The total mass of gas, $M_d(\text{H})$, in M 51 out to a radius of 13.6 kpc is $9.2 \times 10^9 M_\odot$, including the correction given by the simulations (or $6.6 \times 10^9 M_\odot$ without this correction). For M 83, M_{gas} is $7.2 \times 10^9 M_\odot$ out to a radius of 12.2 kpc. For the adopted parameters (i.e. distance, intermediate- κ_ν , etc.), the observed 1.1 mm flux and the

derived mass imply average dust temperatures of 19.4 K and 13.9 K for M 51 and M 83, respectively. For both galaxies,

$$\frac{M_d(\text{H})}{\nu L_\nu(1.1 \text{ mm})} = (3.6 \pm 0.6) \times 10^2 M_\odot / L_\odot \quad (1)$$

where the uncertainty represents the range of values for this sample of two. Using the abovementioned dust temperatures and assuming $\beta = 2.0$, the $M_d(\text{H})/\nu L_\nu(500 \mu\text{m})$ is roughly $29 \pm 8 M_\odot / L_\odot$. This is consistent with Groves et al. (2015), who found that ratio to be 20–30 M_\odot / L_\odot (as inferred from their table 8 for massive galaxies).

The continuum derived and spectral line derived masses and a breakdown of gas masses in M 51 and M 83 as well as a comparison between the two are given in Table 3.

4.3 X-factor

Maps of the X-factor, X_F , are computed from the data, the details of which are given in Appendix C. The mean values of the X-factor for each galaxy are given in Table 4. The X-factor maps radial profiles

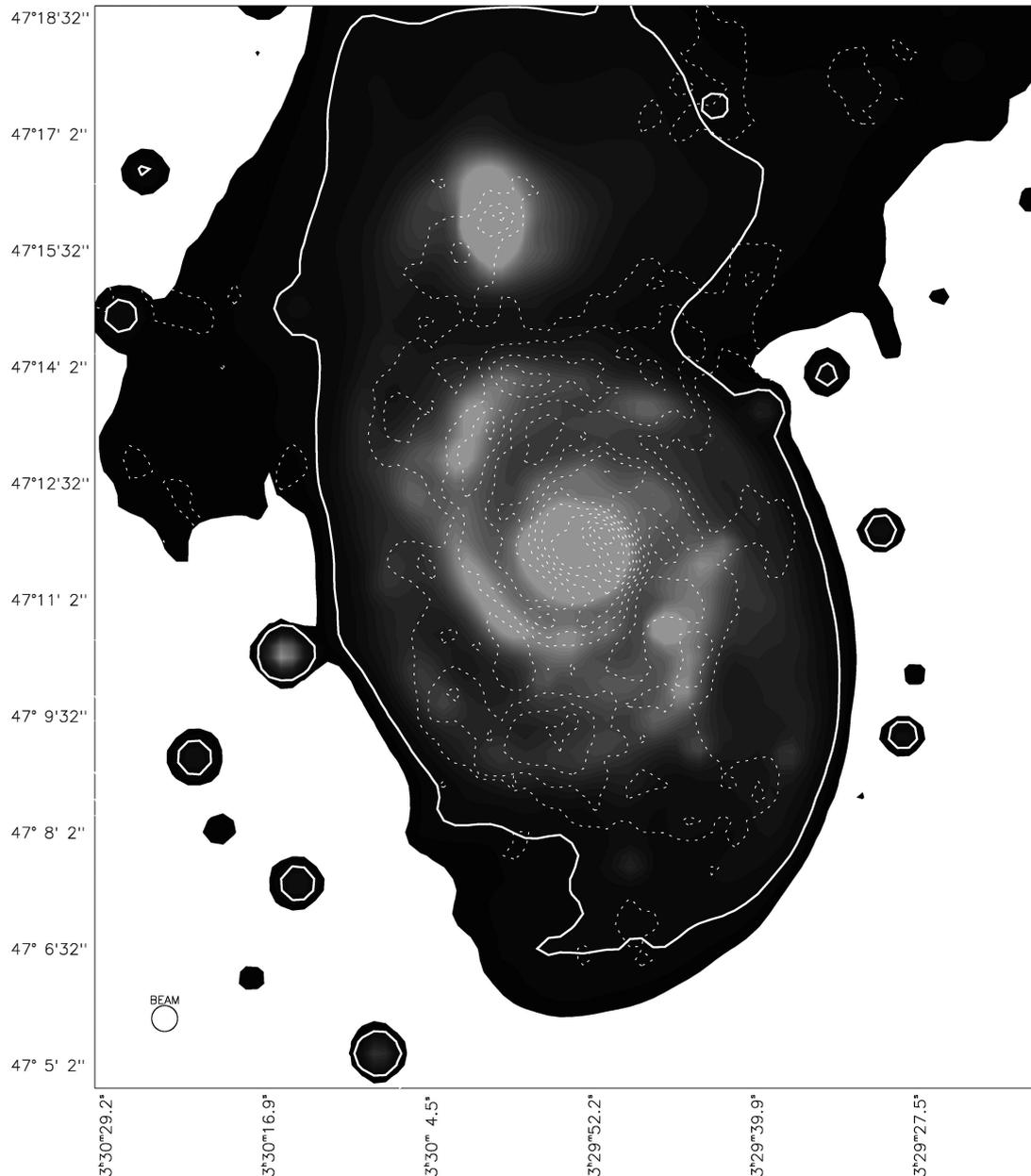


Figure 3. M 51 continuum map at wavelength 1.1 mm (white, dashed contours) is superposed on an image at visible wavelengths (Sloan *g* band; see Baillard et al. 2011) that has been smoothed to the resolution of the AzTEC observations. The solid white contour represents the 25 mag arcsec² isophote. The effective beam is shown in the lower-left corner.

of Figs 12, 13, 14, and 15 reveal spatial variations within the inner 7 kpc radius of both galaxies. The X_F in M 51 is on average larger in the interarm regions than in the arms by factors of roughly 2 or more. This is less obvious in M 83 where the spatial resolution is lower (55 arcsec due to the CO map), but is also roughly the case.

The radial profiles of X_F as seen in Figs 13 and 15 show that X_F does not vary radially by more than a factor of 2 within the central 7 kpc radius. In M 51, such variations are consistent with a constant value for the inner 7 kpc radius to within the errors. In contrast, X_F varies significantly within the central 3 kpc radius of M 83: the ratio of the X_F at the bar ends to that in the central ‘plateau’ is 0.50 ± 0.04 , significantly different from unity. Even in the extreme low- and high- κ_v cases, these results are little changed.

The X -factor maps of both M 51 and M 83 hint at spiral structure. Applying the Fourier spiral analysis mentioned in Section 4.5 to

these X -factor maps yields a two-armed spiral structure with a phase shift of roughly 90° with respect to the main spiral arms for both galaxies, meaning that the interarm/arm X_F ratios for M 51 and M 83 are greater than unity with values 2.5 and 1.5, respectively. These numbers apply only for the inner 8 kpc and inner 6 kpc radii for M 51 and M 83, respectively. The uncertainties of these ratios is about 35 per cent for M 51 and 10 per cent for M 83. In the extreme low- and high- κ_v cases, the inferred interarm/arm ratios are the same to within the uncertainties.

The CO $J = 1 \rightarrow 0$ maps are noisy beyond galactocentric radii of about 7 to 8 kpc. Consequently, we are only able to compute rough lower limits for X_F in the outer discs of M 51 and M 83. (See Appendix C for details.) For M 51, even in the high- κ_v (low- κ_v) case, the lower limits to X_F are roughly $1\text{--}30 X_{20}$ ($7\text{--}10^3 X_{20}$) depending on the radius in the outer disc. For M 83, even in the

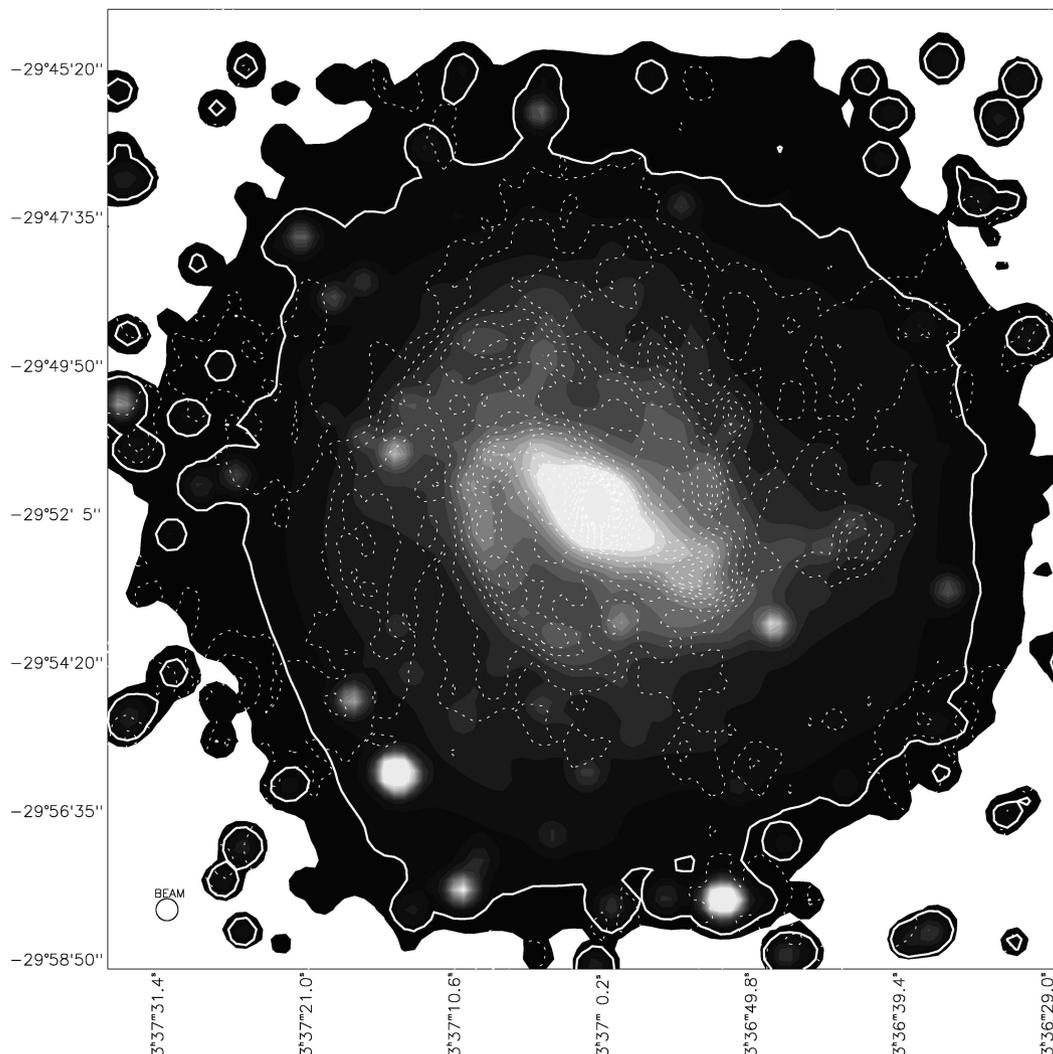


Figure 4. M83 continuum map at wavelength 1.1 mm (white, dashed contours) is superposed on an image at visible wavelengths (R band; see Meurer et al. 2006) that has been smoothed to the resolution of the AzTEC observations. The solid white contour represents the 25 mag arcsec² isophote. The effective beam is shown in the lower-left corner.

high- κ_p (low- κ_p) case, these rough lower limits to X_F are $0.3\text{--}20 X_{20}$ ($2\text{--}600 X_{20}$). This is a strong hint that something unusual might be occurring in the gas or dust (or both) of the outer discs of M 51 and M 83 regardless of the assumed value for the dust mass-absorption coefficient. See Section 5.1 for details.

4.4 Star formation versus gas surface density

Now we examine the variation of the SFR surface density with the gas and dust surface density tracers. (See Appendices D and E for details in producing and comparing such surface density maps.) The plots of the surface densities of SFR versus gas are displayed in Fig. 16. All that figure’s panels, except the upper right, reveal an apparent gas surface density threshold of about $15 M_\odot \text{pc}^{-2}$ at which the SFR surface density rises nearly two orders of magnitude from $\sim 10^{-4}$ to $\sim 10^{-2} M_\odot \text{yr}^{-1} \text{kpc}^{-2}$. Above this threshold, the SFR surface density follows a power-law rise with slopes of 2.5–2.9 for column densities inferred from continuum and with slopes of 1.2–1.6 for column densities inferred from spectral lines.

These superlinear slopes are consistent with the inside-out star formation scenario inferred for spiral galaxies by other means (e.g.

González Delgado et al. 2014). In addition, the surface density threshold of $\sim 15 M_\odot \text{pc}^{-2}$ visible in most of the panels is roughly consistent with the threshold of $\sim 10 M_\odot \text{pc}^{-2}$ determined from the simulations of Clark & Glover (2014). The simulations of Dobbs (2015) suggest a similar threshold in their fig. 11.

Bigiel et al. (2008) also examined the star-formation law in external galaxies. Their fit using $H\alpha$ line and $24 \mu\text{m}$ data had a slope of 1.18. This is comparable to the slope of 1.23 ± 0.01 for the corresponding plot of the current work. The current work shows us that the slope changes yet again – to 2.47 ± 0.05 – when using continuum emission and the H_1 line as a tracer of the molecular gas, i.e. Σ_{SFR} versus $[\Sigma_{\text{d}}(\text{gas}) - \Sigma_{\text{H}_1}]$. A similar difference in slopes between that for Σ_{SFR} versus Σ_{H_2} and that for Σ_{SFR} versus $[\Sigma_{\text{d}}(\text{gas}) - \Sigma_{\text{H}_1}]$ is seen for M 83. These varying numerical values of the slope and how they may or may not affect the physical interpretation of the relationship between the star formation and the gas will be discussed further in Section 5.2.

More insights into star formation are provided by normalizing the SFR surface density to that of the molecular gas, producing what is often called the star formation efficiency (SFE). Images of the Σ_{SFR} divided by molecular gas surface density tracers, either

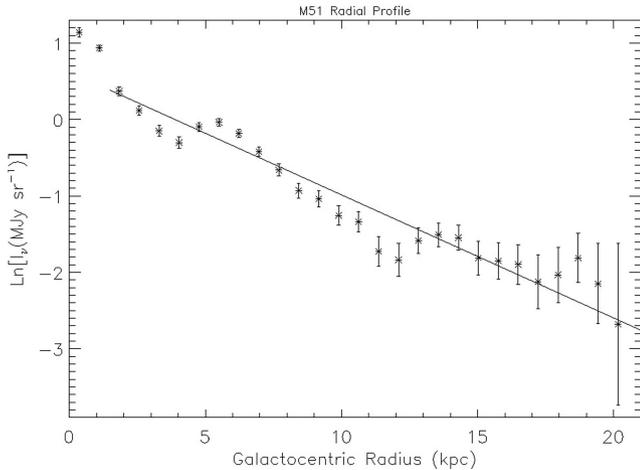


Figure 5. The M51 radial profile in the continuum at wavelength 1.1 mm. The natural logarithm of the 1.1 mm surface brightness of the azimuthally averaged image is plotted against the galactocentric radius in kiloparsecs. The averages are determined within concentric annuli where each annulus is 2 pixels wide (18 arcsec or 0.73 kpc). The solid line is a linear fit to the logarithms of the surface brightnesses for radii outside of the central region (see the text for details).

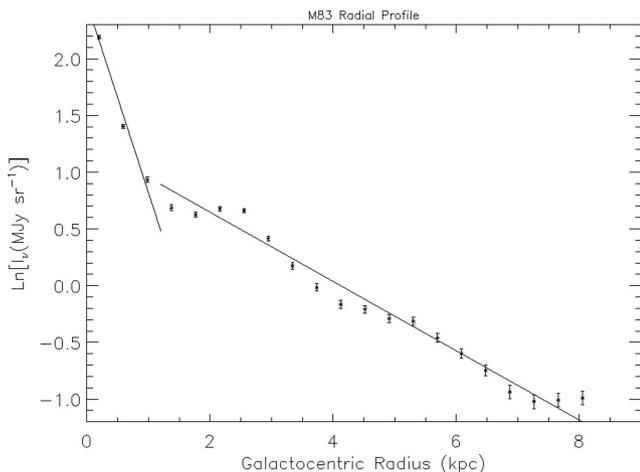


Figure 6. The M83 radial profile in the continuum at wavelength 1.1 mm. The natural logarithm of the 1.1 mm surface brightness of the azimuthally averaged image is plotted against the galactocentric radius in kiloparsecs. The averages are determined within concentric annuli where each annulus is 2 pixels wide (18 arcsec or 0.39 kpc). The solid lines are linear fits to the logarithms of the surface brightnesses for radii inside and outside of the central region (see the text for details).

Table 2. Exponential disk scalelengths for different tracers ^a.

Tracer	M51	M83
Near-IR	$1.8 \pm 0.2^{b,c}$	$2.4 \pm 0.4^{d,e}$
CO $J = 1 \rightarrow 0$	2.8 ± 0.1^c	2.9 ± 0.3^e
1.1 mm	7.3 ± 0.8^f	3.3 ± 0.3^f

^aAll scalelengths in kiloparsecs for adopted distances given in Table 1.

^b3.6 μm .

^cRegan (2006).

^d2.2 μm .

^eLundgren et al. (2004).

^fExcluding central 1.5 kpc radius for M51 and 1.2 kpc radius for M83.

($\Sigma_{\text{d(gas)}} - \Sigma_{\text{H}_1}$) or $2 X_{\text{F}}I(\text{CO})$, are given in Figs 17 and 18. The yellow area surrounding much of the $\Sigma_{\text{SFR}}/\Sigma_{\text{H}_2}$ image of M51 in the upper-right panel suggests a very high SFR per unit gas mass in M51's outer disc. But this is nothing more than an artefact due to using a constant and artificially low X_{F} rather than a higher, and likely more realistic, X_{F} for this outer area of the disc. The top image in fig. 5 of Foyle et al. (2010) is that of their SFE of M51 and the same artefact appears in the form of the white patches on the outer edges of the image. (These white patches also appear in the outer edges of the images of the other two galaxies, NGC 628 and NGC 6946, in that figure.)

Whether or not spiral structure is visible in Figs 17 and 18 is important for determining whether the spiral arms enhance the SFR beyond the corresponding enhancement in the gas surface density due to the arms. Applying spiral arm decomposition to the images using the dust-continuum derived surface densities yields arm/interarm ratios of 1.3 and 2.0 for M51 and M83, respectively. For the M51 image using CO as the molecular gas tracer (right-hand panel), the SFR normalized to the gas surface density is higher between the arms with interarm/arm ratio of 2.4. For M83, this is 1.0. These ratios suggest some mild effect on the SFR due to the spiral arms, but with the SFR normalized to surface densities from the continuum tracer suggesting an opposite effect to that suggested by the SFR normalized to the surface densities from CO. These apparently opposing effects are reconcilable when one considers the spatial variations of $\tau_{\text{vd}}/N(\text{H})$ and of X_{F} as discussed in Section 5.2.

4.5 Spiral structure Fourier analysis

Given that the large-scale structure of M51 and M83 has been reliably recovered according to the simulations, it is worthwhile to examine the spiral structure of both galaxies. Hence, a 2D Fourier analysis on the basis of logarithmic spirals (Kalnajs 1975; Considère & Athanassoula 1988; Puerari & Dottori 1992; Block & Puerari 1999) was conducted on different images of both M51 and M83, thereby allowing tests of the spiral structure.

One obvious test is to see whether the spiral structure observed in visible (or nearly visible) light, due to stars, is the same as that observed in the millimetre continuum, due to dust (and its associated gas). Accordingly, the spiral-arm analysis mentioned above is applied to an *R*-band image of M51¹ and an *I*-band image of M83,² as well as the AzTEC 1.1-mm image. Fig. 19, for example, shows that the power spectrum in M51 for the $m = 2$ (i.e. two-arm) spiral pattern is identical, to within the uncertainties, between the *R* band and the 1.1 mm continuum M51. However, the power spectrum in Fig. 20 for M83 reveals a noticeable difference between the spiral structure in the 1.1 mm image and that in the *I*-band image; only a hint of a small bar is discernible in the former image, whereas a more prominent bar adorns the latter image. Further analysis finds that the visible (or near visible) light and 1 mm continuum spirals are not offset from each other. Accordingly, this suggests that the stars and gas are also not offset from each other in the spiral arms.

Other tests of spiral structure are applied to the X-factor and to the SFR surface density normalized to the gas surface density, $\Sigma_{\text{SFR}}/\Sigma_{\text{gas}}$ or $\Sigma_{\text{SFR}}/\Sigma_{\text{H}_2}$. The results of these tests are presented in Figs 21 and 22 for M51 and M83, respectively, which display

¹ See http://ned.ipac.caltech.edu/cgi-bin/ex_refcode?refcode=1994DSS...1...0000percent3AforM51.

² See http://ned.ipac.caltech.edu/cgi-bin/ex_refcode?refcode=2000ApJS...131..441KforM83.

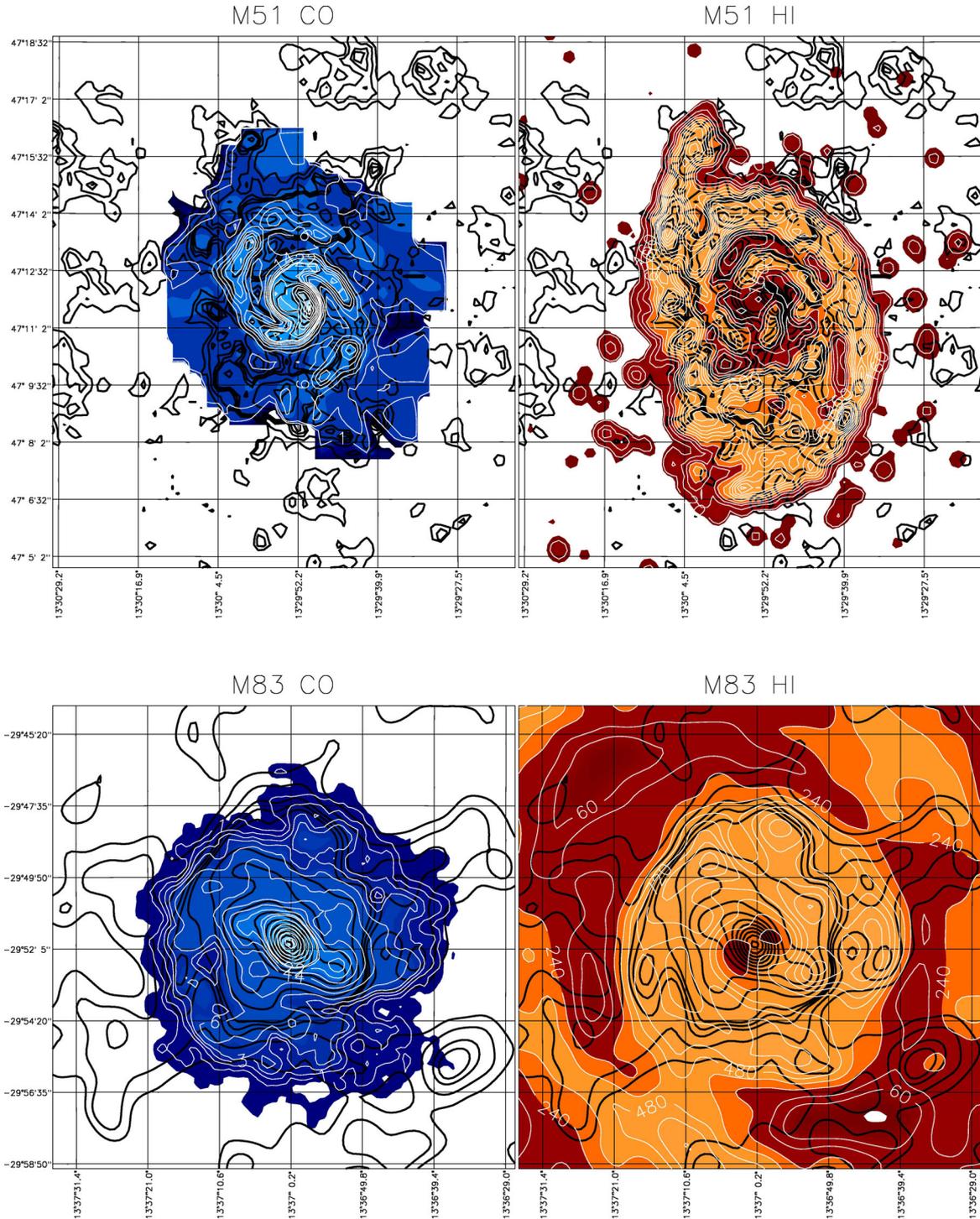


Figure 7. CO $J = 1 \rightarrow 0$ and H I maps are superposed on the corresponding 1.1 mm continuum maps of M 51 and M 83. The 1.1 mm continuum map in each panel is represented by the dark contours with levels of 0.4, 0.6, 0.8, 0.9, 1.0, 1.4, 1.8, 2.2, ..., 3.8 MJy sr^{-1} for M 51 with the same sequence of levels for M 83, but with a maximum level of 5.0 MJy sr^{-1} . The 1.1 mm spatial resolution is 20 arcsec in the M 51 panels and 55 arcsec in the M 83 panels. In the M 51 CO (upper left) panel, the white contours and blue shading show the CO $J = 1 \rightarrow 0$ emission convolved to a 20 arcsec resolution with contour levels of 1, 2, 3, 4, 6, 10, 14, 18, 22, ..., 76 K km s^{-1} . In the M 51 H I panel, the white contours and red/orange shading show the H I emission convolved to a 20 arcsec resolution with contour levels of 20, 80, 160, 240, 360, 480, ..., 1560 K km s^{-1} . The M 83 CO panel displays the CO $J = 1 \rightarrow 0$ image at a 55 arcsec resolution with the same sequence of contour levels as for the M 51 image, but with 68 K km s^{-1} as the maximum level. The M 83 H I panel depicts the H I emission convolved to 55 arcsec resolution with white contours and red/orange shading; the contour levels are 60, 120, 240, 360, 480, ..., 1320 K km s^{-1} . The lowest contour level is roughly equivalent to 1σ for all maps in this figure.

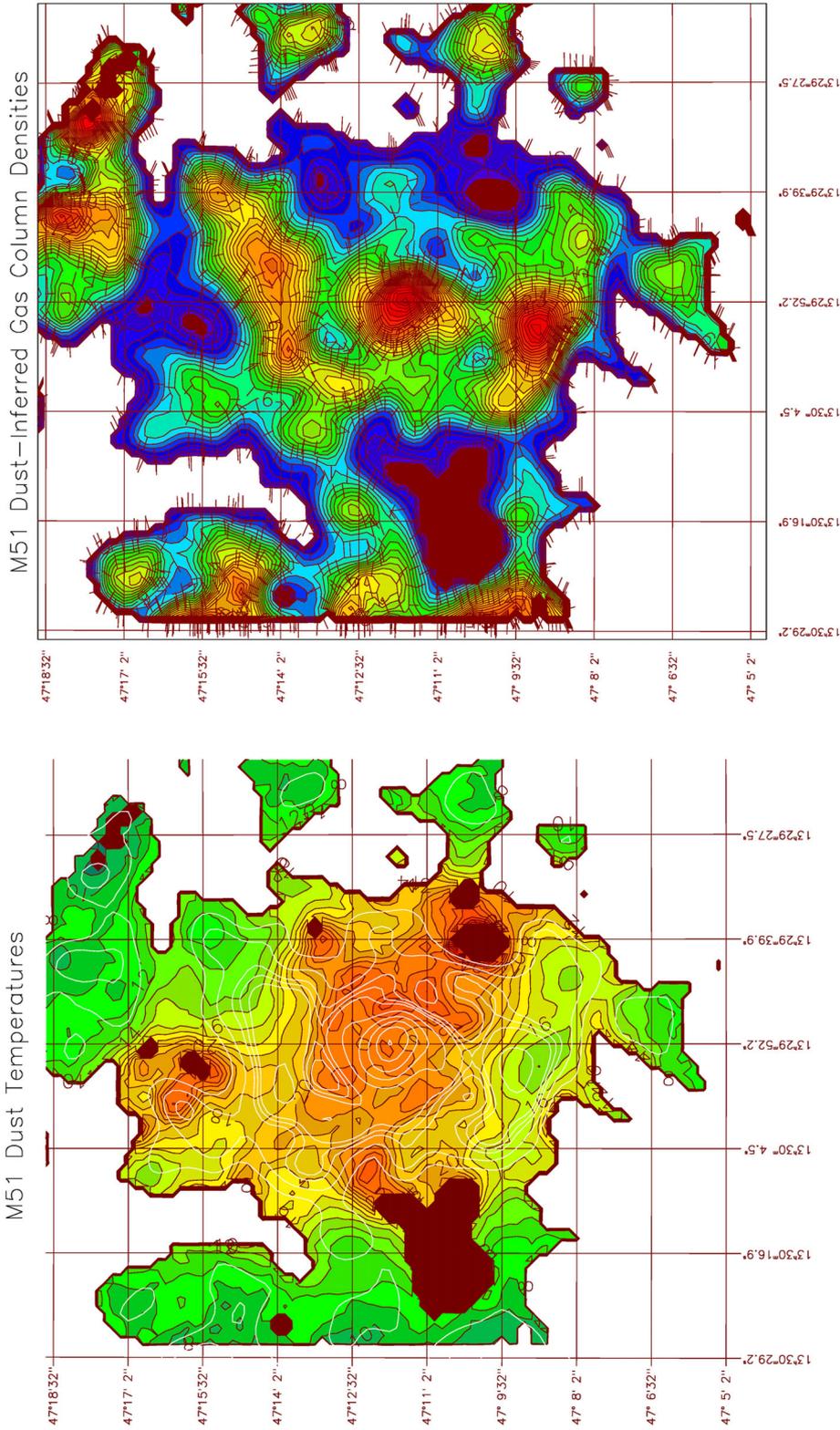


Figure 8. M51 dust temperature map and H — nuclei column density map are shown in the left- and right-hand panels, respectively. The coordinates are epoch 2000.0. Left-hand panel: the dust temperatures are those determined from the ratio of the *Spitzer* 156 μm to AzTEC 1.1 mm intensity ratio. The coloured contours give the dust temperatures in 1 K steps from 8 K. The dark solid areas represent regions where the signal to noise of the intensity ratio was less than unity. The white contours give the 1.1 mm intensities in MJy sr^{-1} for the 1.1 mm continuum map degraded to the 38 arcsec resolution of the 156 μm map. The 1.1 mm contour levels are 0.4, 0.6, 0.8, 0.9, 1.0, 1.4, 1.8, 2.2, 2.6, 3.0 MJy sr^{-1} . Right-hand panel: these are gas column densities inferred from the dust-continuum emission (see the text). The contour levels are 0.5, 1.0, 1.5, 2.0, \dots , $20.0 \times 10^{20} \text{H} — \text{nuclei cm}^{-2}$. The tick marks point towards lower contour levels. The dark solid areas correspond to those of the dust temperature map in the left-hand panel.

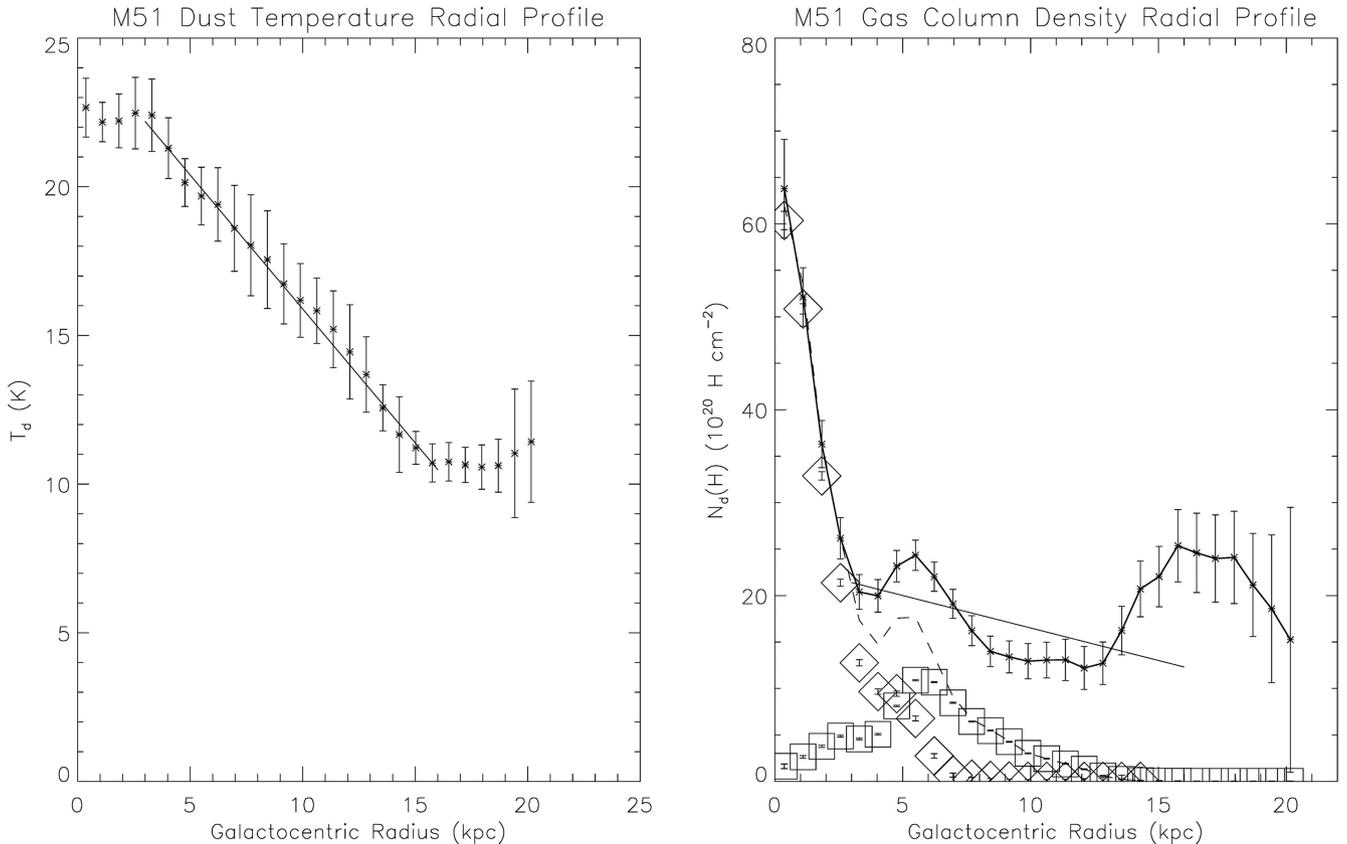


Figure 9. The M 51 radial profile of the $156\ \mu\text{m}/1.1\ \text{mm}$ dust temperature and of the H — nuclei column densities, $N_d(\text{H})$, as inferred from the dust continuum emission are shown in the left- and right-hand panels, respectively. The azimuthally averaged dust temperature and column densities are plotted against the galactocentric radius in kiloparsecs. The averages are determined within concentric annuli where each annulus is 2 pixels wide (18 arcsec or 0.73 kpc). The solid line in the left-hand panel is a linear fit to radii from 3 to 16 kpc, where the dust temperature has a linear decline of about $0.9\ \text{K}\ \text{kpc}^{-1}$. The radial profile of the H — nuclei column densities is given in the right-hand panel above by the thick solid curve joining the data points. The thin solid line in the right-hand panel is a linear fit to the column densities at radii from 3 to 16 kpc, which corresponds to where the dust temperature has a linear decline. The diamonds represent the column densities of H — nuclei in the molecular gas as inferred from CO $J = 1 \rightarrow 0$ using a constant X-factor (see Section 4.3). The squares represent the column densities of H — nuclei in the atomic gas as inferred from H I 21-cm emission. For both the CO and H I, the error bars are smaller than the symbols. The dashed line gives the total gas column density as inferred from both CO and H I. Note that the additive correction of $5.5 \times 10^{20}\ \text{H — nuclei}\ \text{cm}^{-2}$ was not applied to the $N_d(\text{H})$ points above.

the results of this Fourier analysis for the 1.1 mm continuum images, the X-factor maps, and the maps of the SFR normalized to the molecular gas surface density, where that surface density is determined from the continuum and H I for the ‘SFA’ panel and determined from CO for the ‘SFAG’ panel. Clearly the higher values for both the X-factor, and the SFAG images are found in the interarm region, and the spiral structure have almost the same pitch angle as the main AzTEC arms. For the SFA images, the pitch angle is a bit smaller, but it is in phase with the arms we detect in the AzTEC image. These results are similar between M 51 and M 83.

We have used the detected positions of the arms in the AzTEC images of M 51 and M 83, and calculated the arm-to-interarm ratios of the processed images. These are presented in Table 5. The arm-to-interarm ratios are for the inner discs of both galaxies – galactocentric radii of 1.8–5.5 kpc and 1.0–2.9 kpc for M 51 and M 83, respectively. These radii were chosen for consistency with the X-factor maps. Most of the numerical values listed in Table 5 are within a factor of 2 of unity. The spiral arms seen in the 1.1 mm continuum have ratio values that are comparable to those seen in red light or I band. The arm/interarm values of the X-factor images indicate that the X-factor is higher in the interarm regions,

where those interarm regions are defined as those in the millimetre continuum and in the I and R bands (i.e. between the dust and stellar arms).

Similar to the X-factor, the SFAG map is higher between the arms than in the interarms, at least for M 51. Given that the SFAG map was computed from the X-factor, it is not surprising that both X-factor and SFAG maps have this quirk in the arm-to-interarm ratio. Compensating the arm/interarm ratio of SFAG for that of the X-factor suggests that the ‘true’ arm-to-interarm ratio is above unity. Indeed, that is confirmed in the SFA map, which is the SFR surface brightness normalized to the molecular gas surface density determined from the continuum and the H I line and, as a result, the SFA map is independent of the X-factor.

Taken at face value, this suggests that star formation is more ‘efficient’ in the arms than in the interarms, in the sense that the SFR in the arms is enhanced beyond that expected from simply having more gas and dust surface density in the arms. However, the values of the arm-to-interarm ratios presented here are subject to systematic effects. See Section 5.2 for more discussion of this. In any event, even if real, this enhancement is less than a factor of 2 or even 1.5.

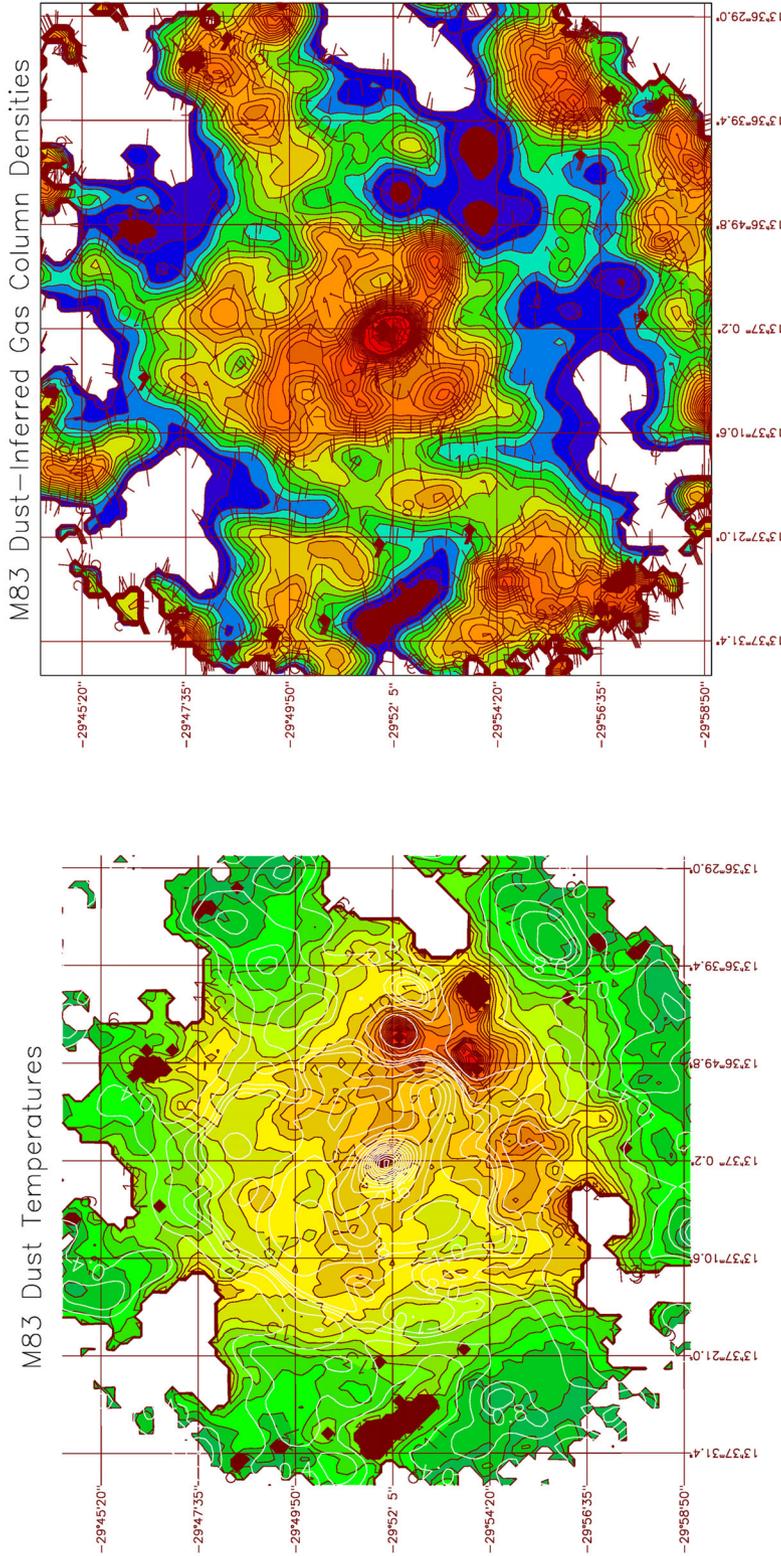


Figure 10. M 83 dust temperature map and H — nuclei column density map are shown in the left- and right-hand panels, respectively. The coordinates are epoch 2000.0. Left-hand panel: the dust temperatures are those determined from the ratio of the *Spitzer* 156 μm to AZTEC 1.1 mm intensity ratio. The coloured contours give the dust temperatures in 1 K steps from 9 K. The dark solid areas represent regions where the signal to noise of the intensity ratio was less than unity. The white contours give the 1.1 mm intensities in Mjy sr^{-1} for the 1.1 mm continuum map degraded to the 38 arcsec resolution of the 156 μm map. The 1.1 mm contour levels are 0.4, 0.6, 0.8, 0.9, 1.0, 1.4, 1.8, 2.2, ..., 7.0 Mjy sr^{-1} . Right-hand panel: these are gas column densities inferred from the dust-continuum emission (see the text). The contour levels are 1, 2, 3, 4, ..., 30, 35, 40, 45, $50 \times 10^{20} \text{H — nuclei cm}^{-2}$. The tick marks point towards lower contour levels. The dark solid areas correspond to those of the dust temperature map in the left-hand panel.

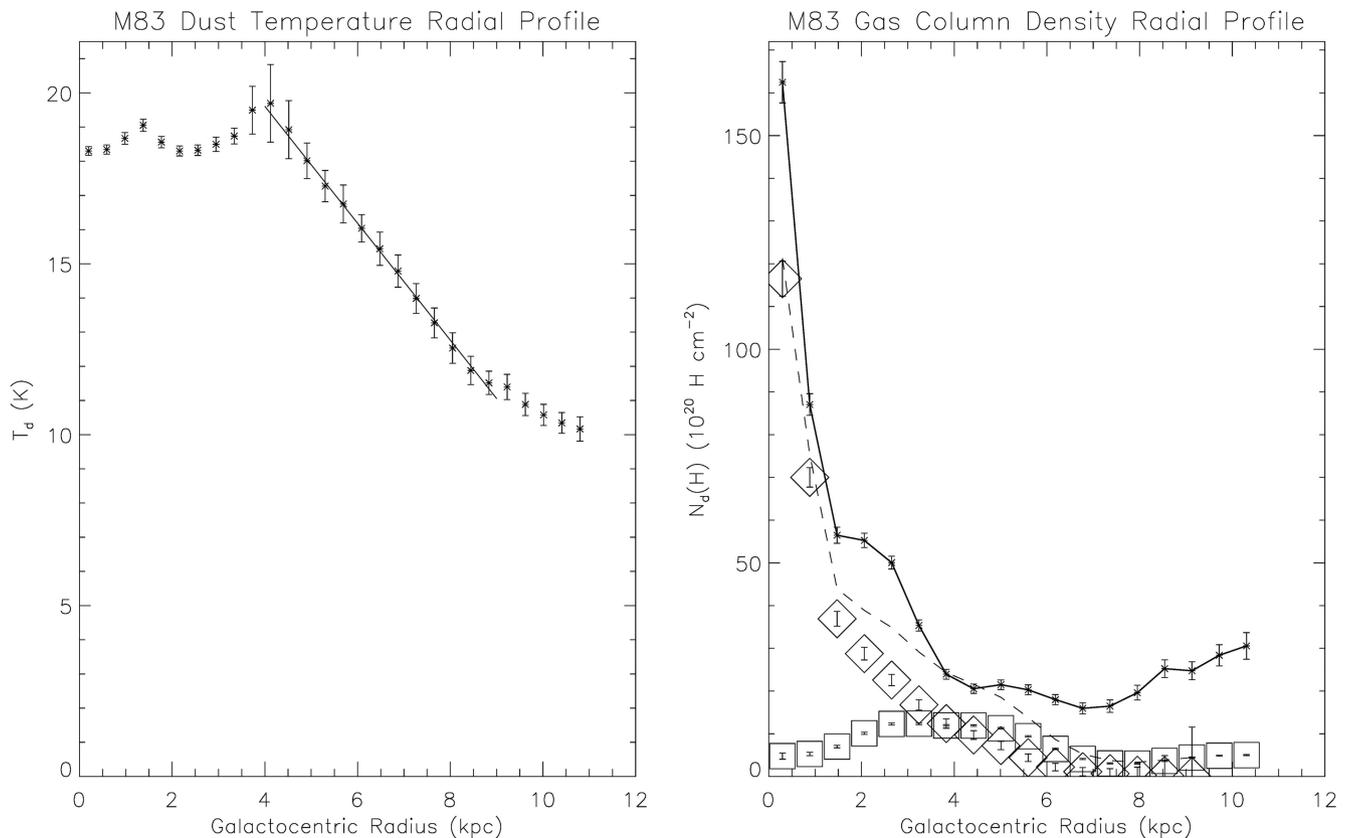


Figure 11. The M 83 radial profile of the $156\ \mu\text{m}/1.1\ \text{mm}$ dust temperature and of the H — nuclei column densities, $N_d(\text{H})$, as inferred from the dust continuum emission are shown in the left- and right-hand panels, respectively. The azimuthally averaged dust temperature is plotted against the galactocentric radius in kiloparsecs. The averages are determined within concentric annuli where each annulus is 2 pixels wide (18 arcsec or 0.39 kpc). The solid line in the left-hand panel is a linear fit to radii from 4 to 9 kpc, where the dust temperature has a linear decline of about $1.7\ \text{K}\ \text{kpc}^{-1}$. The radial profile of the H — nuclei column densities, $N_d(\text{H})$, is given in the right-hand panel and is represented by the thick solid curve joining the data points. The diamonds represent the column densities of H — nuclei in the molecular gas as inferred from CO $J = 1 \rightarrow 0$ using a constant X-factor (see Section 4.3). The squares represent the column densities of H — nuclei in the atomic gas as inferred from H I 21-cm emission. For both the CO and H I, the error bars are comparable to or smaller than the symbols. The dashed line gives the total gas column density as inferred from both CO and H I.

Table 3. Gas masses^a in M 51 and M 83.

	M 51	M 83
$M_d(\text{H})$ ^b	9.2×10^9	7.2×10^9
$M(\text{H}_2)$ < italic > c < /italic >	1.7×10^9 ^d	1.4×10^9 ^e
$M(\text{H I})$	1.8×10^9	2.2×10^9
$M(\text{H}_2) + M(\text{H I})$	3.5×10^9	3.6×10^9
$\frac{M(\text{H}_2) + M(\text{H I})}{M_d(\text{H})}$	0.4	0.5

^aAll masses in M_\odot .

^bTotal gas mass as inferred from dust continuum emission.

^cMolecular gas mass as inferred from the CO $J = 1 \rightarrow 0$ line.

^dUsing $X_F = 0.8 X_{20}$ for M 51, see Section 4.3.

^eUsing $X_F = 1.0 X_{20}$ for M 83, see Section 4.3.

5 DISCUSSION

Maps of M 51 and M 83 were made in the 1.1 mm continuum from observations with the instrument AzTEC with the JCMT. Combining with these maps with the corresponding *Spitzer* 160 μm (or, more properly, 155.9 μm) maps gave estimates of the gas surface densities in these two galaxies (see Appendices B and B1 for a detailed discussion). With these gas surface density maps, spatial variations of the X-factor were estimated. In addition, we investi-

Table 4. Mean^a X-factorestimates^b in M 51 and M 83.

M 51	M 83
0.8 ^c	1.0

^aThe $1/\sigma^2$ -weighted means of the X_F map and only for the inner 7 kpc radius. Uncertainty of about ± 50 per cent.

^bIn units of X_{20} or $10^{20}\ \text{H}_2\ \text{cm}^{-2}\ (\text{K}\ \text{km}\ \text{s}^{-1})$.

^cIncludes the correction determined from the simulations.

gated the relationship between the gas surface density and that of the SFR. These are dealt with in more detail below.

5.1 The X-factor, its spatial variations, and CO-dark gas

The most important results of this work regarding the X-factor are the following.

- (i) The average X-factor for each galaxy can be estimated from the current observations, even if crudely. Those average values are ~ 0.8 and $\sim 1.0 X_{20}$ for M 51 and M 83, respectively.
- (ii) The X-factor is higher in the interarm regions than in the arms.
- (iii) There seems to be CO-dark gas that resides mostly in the outer discs of both M 51 and M 83.

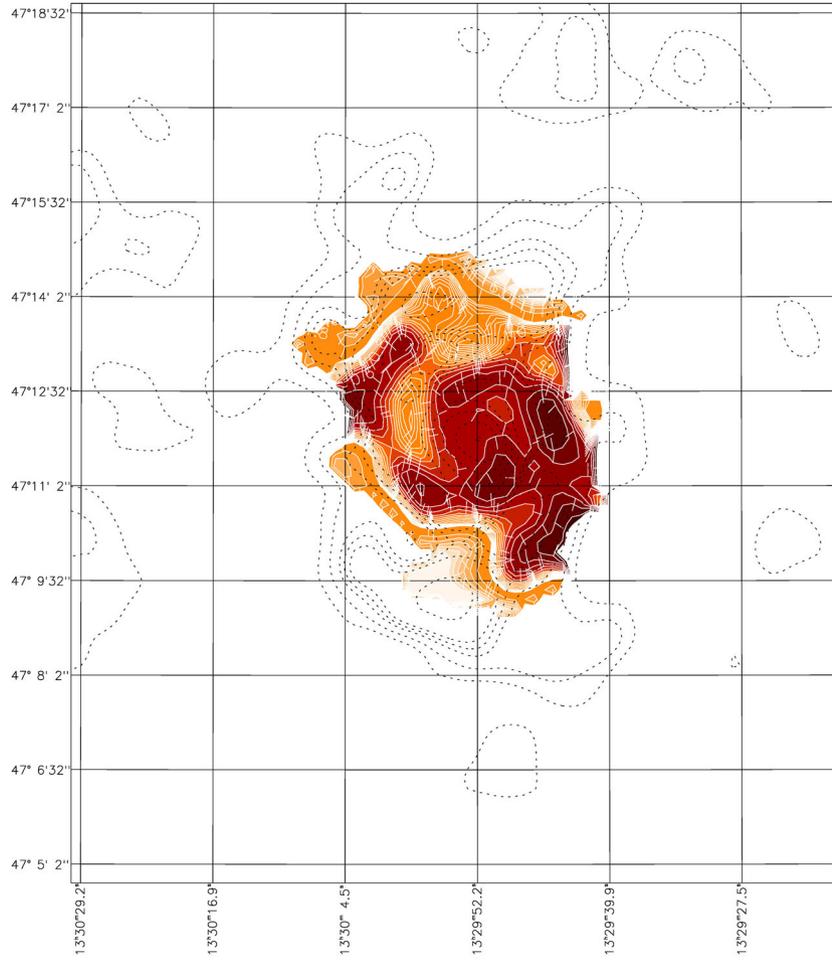


Figure 12. M51 X_F map is shown with contour levels 0.2, 0.3, 0.4, ..., 2.0, 4.0, 6.0, ..., 14.0 X_{20} . The dotted contours are the 1.1 mm surface brightness with levels of 0.4, 0.6, 0.8, 0.9, 1.0, 1.4, 1.8, ..., 3.0 MJy sr^{-1} .

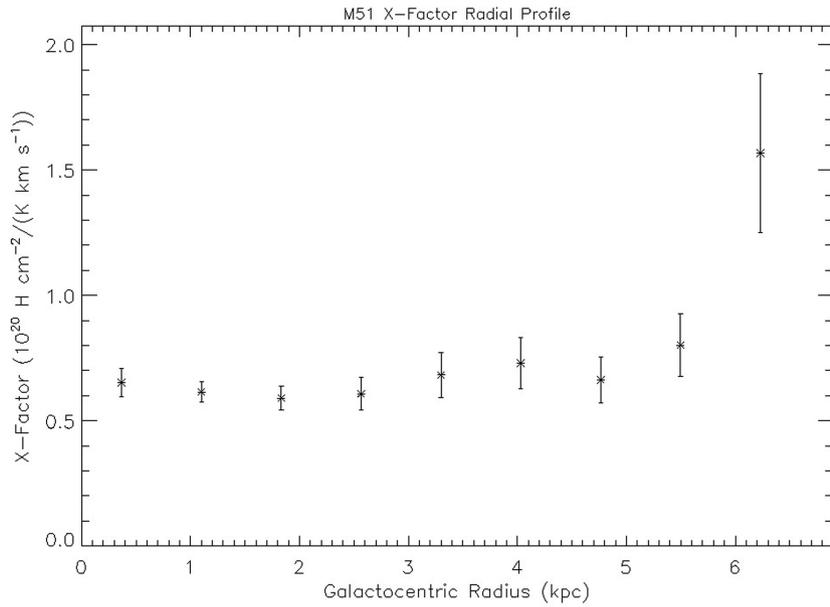


Figure 13. The M51 radial profile of the X-factor. The azimuthally averaged X-factor is plotted against the galactocentric radius in kiloparsecs. The averages are determined within concentric annuli where each annulus is 2 pixels wide (18 arcsec or 0.73 kpc). Note that the 23 per cent upward correction to the above X-factor values determined from the simulations has not been applied.

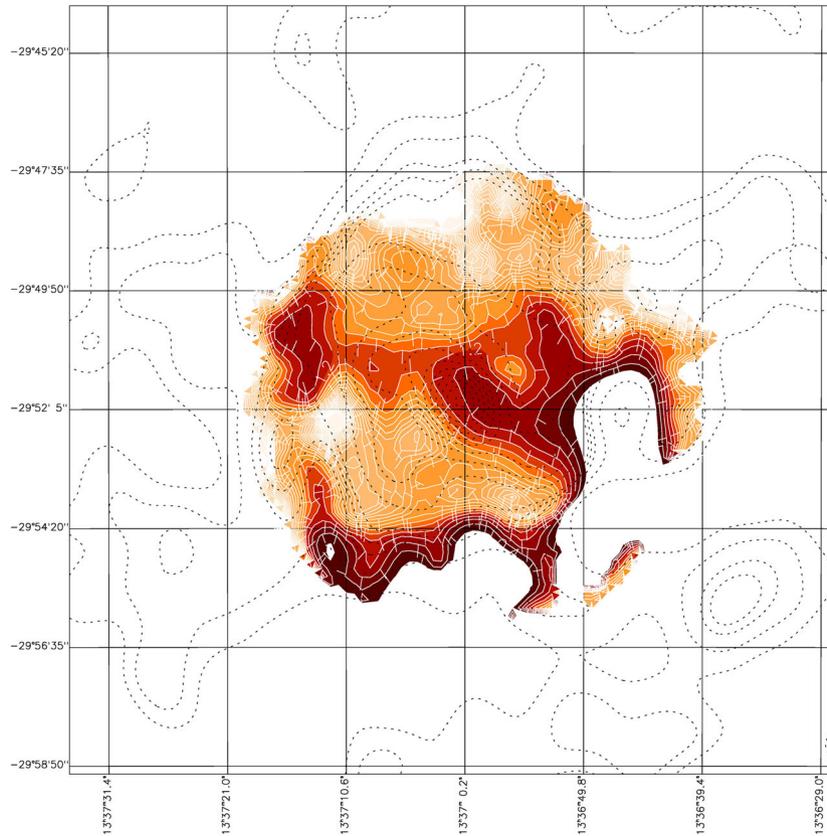


Figure 14. M83 X_F map is shown with contour levels 0.4, 0.6, 0.8, ..., 5.2 X_{20} . The dotted contours are the 1.1 mm surface brightness with levels of 0.4, 0.6, 0.8, 0.9, 1.0, 1.4, 1.8, ..., 5.0 MJy sr^{-1} .

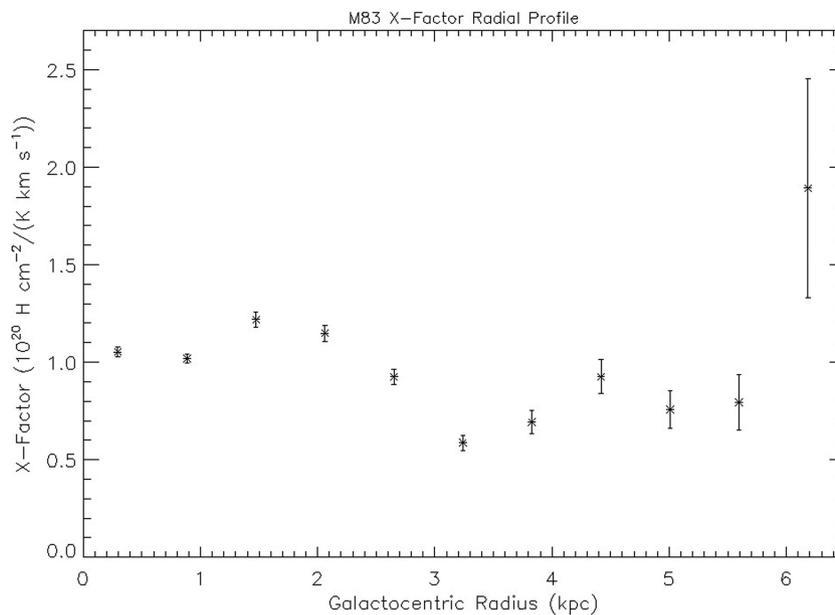


Figure 15. The M83 radial profile of the X-factor. The azimuthally averaged X-factor is plotted against the galactocentric radius in kiloparsecs. The averages are determined within concentric annuli where each annulus is 3 pixels wide (27 arcsec or 0.59 kpc).

The latter two results are robust to a range of adopted $\tau_{\text{vd}}/N(\text{H})$ values.

The variation of the X-factor spatially and from source to source could be due, in part, to variations in metallicity. Theoretical work using the observational data also support a dependence of X_F on

metallicity (e.g. Lagos et al. 2012; Narayanan et al. 2012). In contrast, Sandstrom et al. (2013) do not find a strong correlation of X_F with metallicity. However, their sample only had a metallicity range of 0.5–0.8 dex within factors of 3 of solar. The irregular galaxies observed by Israel and others (e.g. see Israel 1988;

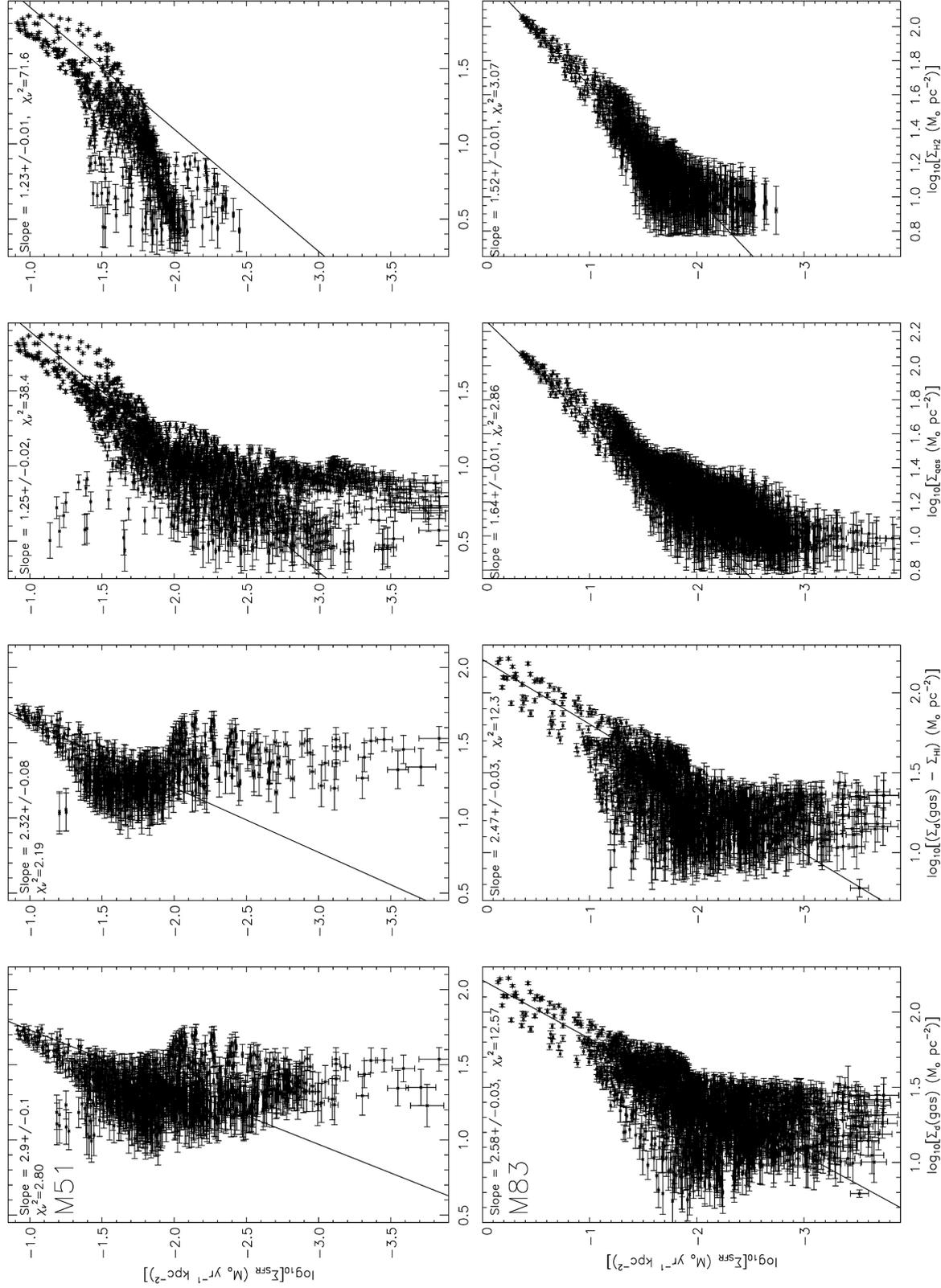


Figure 16. The logarithm of the star formation rate (SFR) surface density is plotted against that from a tracer of gas surface density in each of the four panels for M51 (upper) and for M83 (lower). In the first two panels for each row, the gas surface densities are determined from the dust continuum. In the last two panels for each row, these surface densities are determined from gas spectral lines only. The column densities in the first and third panels of each row are from tracers of the total gas (molecular + atomic) surface density and those in the second and fourth panels are from tracers of molecular gas surface density only. The linear fits are to the points where the $\log_{10}[\Sigma_{\text{SFR}}(\text{M}_{\odot} \text{yr}^{-1} \text{kpc}^{-2})] \geq -2.0$. The value of the fitted slope and the reduced chi-square of the fit appear in the upper left of each panel.

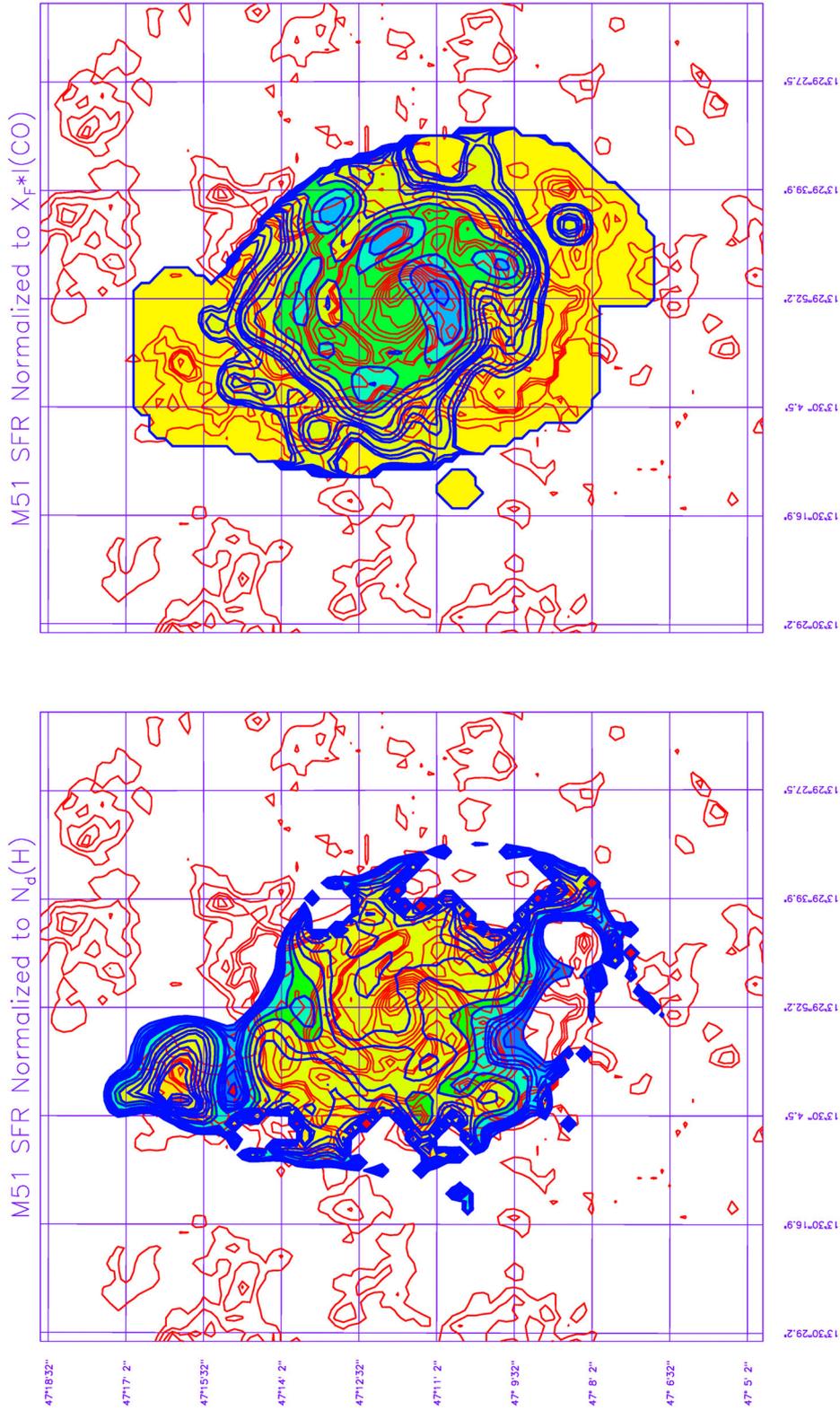


Figure 17. Maps of the SFR normalized to the molecular gas surface density – sometimes called the star formation ‘efficiency’ – are shown for M 51. The left-hand panel has the SFR normalized to the difference between the continuum-derived gas surface density, $\Sigma_d(\text{gas})$, and the surface density of atomic gas, $\Sigma_{\text{H I}}$. The right-hand panel is the SFR normalized to the molecular gas surface density estimated from the adopted X-factor applied to the CO $J = 1 \rightarrow 0$ surface brightness. The contour levels for the blue through green through yellow shaded areas for the left-hand panel are 0.25, 0.30, 0.35, ..., 0.5, 0.75, 1.0, 1.5, 3.0, 4.0, 5.0, 6.25, 12.5, 18.75, 25.0 Gyr^{-1} . The contours of the right-hand panel are those of the left-hand panel scaled by 0.4. The red contours in both panels are the 1.1 mm continuum surface brightnesses for M 51 with levels 0.4, 0.6, 0.8, 0.9, 1.0, 1.4, 1.8, 2.2, 2.6, 3.0, 3.4, 3.8 MJy sr^{-1} .

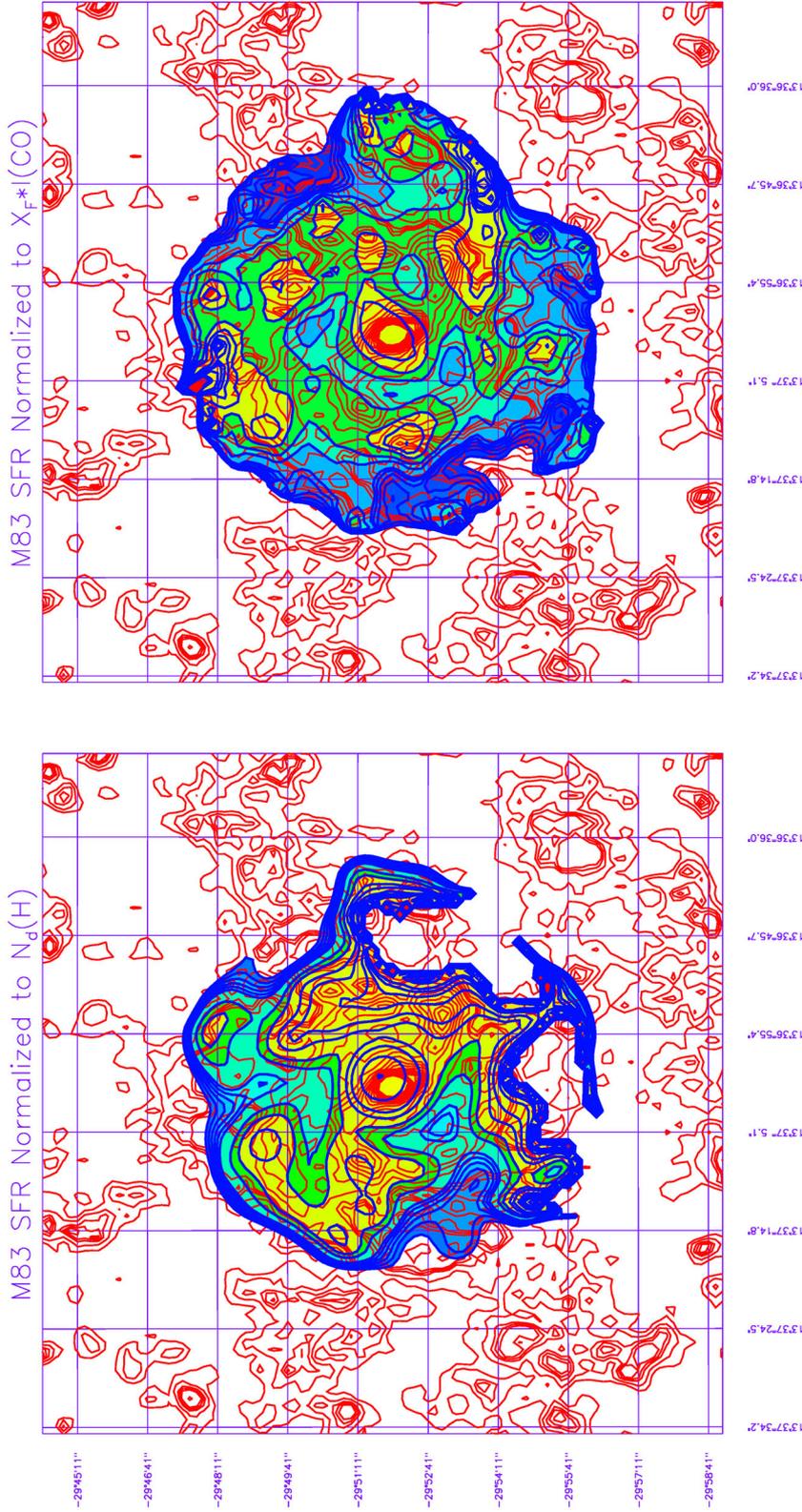


Figure 18. Maps of the SFR normalized to the molecular gas surface density as shown for M 83. The left-hand panel has the SFR normalized to the difference between the continuum-derived gas surface density, $\Sigma_d(\text{gas})$, and the surface density of atomic gas, $\Sigma_{\text{H I}}$. The right-hand panel is the SFR normalized to the molecular gas surface density estimated from the adopted X-factor applied to the CO $J = 1 \rightarrow 0$ surface brightness. The contour levels for the blue through green through yellow shaded areas for the left-hand panel are 0.25, 0.30, 0.35, . . . , 0.5, 0.75, 1.0, 1.5, 3.0, 4.0, 5.0, 9.25, 12.5, 18.75, 25.0 Gyr^{-1} . The contours of the right-hand panels are those of the left-hand panel scaled by 0.4. The red contours in both panels are the 1.1 mm continuum surface brightnesses for M 83 with levels 0.4, 0.6, 0.8, 0.9, 1.0, 1.4, 1.8, 2.2, 2.6, 3.0, 3.4, 3.8, 4.2, 4.6, 5.0, 5.4, 5.8, 6.2, 6.6, 7.0 MJy sr^{-1} .

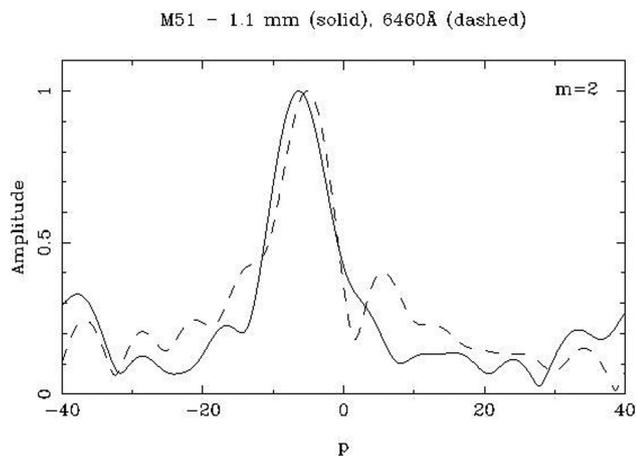


Figure 19. Fourier transform results for the logarithmic spiral in the 1.1 mm and *R*-band images of M51. The power spectrum for the $m = 2$ spiral is shown for both images as a function of p , which is a measure of the pitch angle (i.e. the p is related to pitch angle, P , by $p = -m \cot(P)$). The main peak of each spectrum gives very similar pitch angles for both the 1.1 mm ($P = 18^\circ$) and *R*-band ($P = 20^\circ$) images.

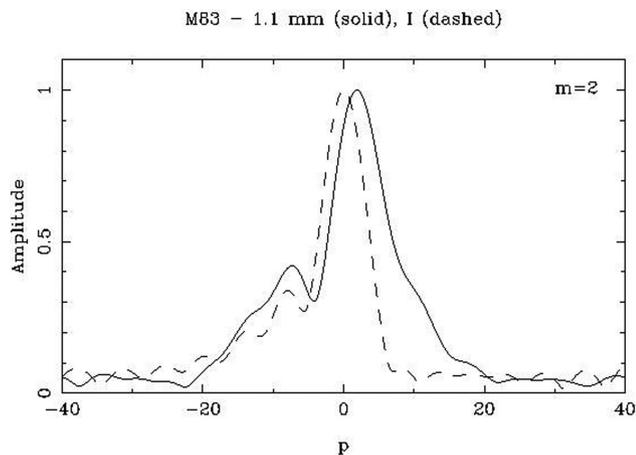


Figure 20. Similar to Fig. 19, Fourier transform results for the logarithmic spiral in the 1.1 mm and *I*-band images of M83. The power spectrum for the $m = 2$ spiral is shown for both images as a function of p . For this galaxy, in the radial range being analysed, the *I* image is dominated by the oval/bar distortion. This power spectrum shows a clear peak at $p = 0$ ($P = 90^\circ$).

Dettmar & Heithausen 1989; Israel 1997a,b; Madden et al. 1997; Fukui et al. 1999) typically had metallicities much less than solar, sometimes only a few per cent of solar, and found X_F -factors an order-of-magnitude or more higher than the standard value. So a strong X_F -metallicity relation may exist for galaxies with strongly subsolar metallicities. For galaxies with roughly solar metallicities, while metallicity alone is apparently insufficient in constraining X_F , it is still relevant. For example, using the data points for NGC4321 from the left-hand panel of fig. 10 in Sandstrom et al. (2013) yields a correlation coefficient of -0.6 . This suggests that each galaxy has its own X_F -metallicity relationship.

Given that subsolar metallicities imply larger X -factors, perhaps the interarm regions of M 51 and M 83 have subsolar metallicities, while being at solar level in their arms.³ The models of Narayanan

³ Parkin et al. (2013), however, find that the FUV flux and molecular gas density are the same for arm and interarm regions in M51. Whether this implies the same metallicity in both regions is unclear.

et al. (2012), as well as their fig. 1, suggest that the metallicity in the interarm regions would be systematically lower by a factor of ~ 2 to 3 in order to increase X_F by a factor of 2 with respect to that in the spiral arms (see their equation 8). While observations of metallicity in the ISM of M 51 and M 83 apparently do not support such a systemically lower metallicity between their spiral arms (see Bresolin, Garnett & Kennicutt 2004; Bresolin et al. 2009), they also do not rule it out: such observations are towards $H II$ regions and are heavily biased towards the spiral arms.

This could, in turn, affect the results of studies of the effects of the spiral arms on star formation. Foyle et al. (2010), for example, looked at whether the SFR normalized to the molecular gas surface density is higher in the spiral arms of three galaxies. They adopted a spatially constant X -factor, which, to within a factor of 2, is likely correct. This will be discussed further in Section 5.2.

The modelling by Narayanan et al. (2012) mentioned above suggests that the X -factor depends mainly on the metallicity of the gas in a galaxy and, to a lesser extent, on the average CO surface brightness. Using the roughly solar metallicities of M 51 and M 83 (see Bresolin et al. 2004, 2009), the observed CO brightnesses from the data used here, and applying expression 8 of Narayanan et al. (2012) yields $X_F \simeq 3 X_{20}$ for M 51 and $X_F \simeq 3-6 X_{20}$ for M 83; this is much higher than the observed values found in the current work. Admittedly, our estimates are uncertain by factors of about 2. Nevertheless, our estimate of $X_F \simeq 0.8 X_{20}$ for M 51 agrees with the result of Nakai & Kuno (1995) who find $X_F = 0.9 \pm 0.1 X_{20}$ from using the observed extinction in $H II$ regions. Accordingly, the theoretical models need further adjustments.

As mentioned previously, there is evidence that adopting a spatially constant X -factor does not account for CO-dark gas. For example, Table 3 suggests that about half the total gas mass is unaccounted for when using the CO $J = 1 \rightarrow 0$ and $H I$ 21-cm spectral lines (see Section 4.3 and Appendix B1). If we adopt a threshold for CO-dark gas that corresponds to an X_F that is a factor of 4 higher than the average for each galaxy, then some of the edges of the shaded regions of Figs 12 and 14 indicate such gas. It is worth noting that these edges are well within the boundaries of the CO maps.

Of course, possible alternative interpretations for the high X -factor or its high lower limits are not entirely ruled out. These include the following.

- (i) An extended low-surface brightness artefact in the 1.1 mm continuum maps of M 51 and M 83.
- (ii) Insufficient mapping of CO $J = 1 \rightarrow 0$ in the outer discs of these galaxies.
- (iii) Dust with unusual properties such as unusually high dust-mass absorption coefficient (i.e. κ_ν) and/or a high dust-to-gas mass ratio.
- (iv) Optically thick $H I$ 21-cm emission.

Each of the above could mimic the presence of CO-dark gas. The first alternative is unlikely given that the simulations of the AzTEC 1.1 mm observations have accurately accounted for any constant offsets in the M 51 and M 83 maps. The second alternative is unlikely because there is evidence for such CO-dark gas seen at radii well within the boundaries of the existing CO maps, as mentioned previously. The third alternative is unlikely because having a combination of high κ_ν and high x_d would not be sufficient for positions with X_F two orders of magnitude larger.

The fourth alternative is a partly valid explanation for CO-dark gas in our Galaxy according to the Planck Collaboration XIX (2011b). They estimated that up to half of the dark gas could be

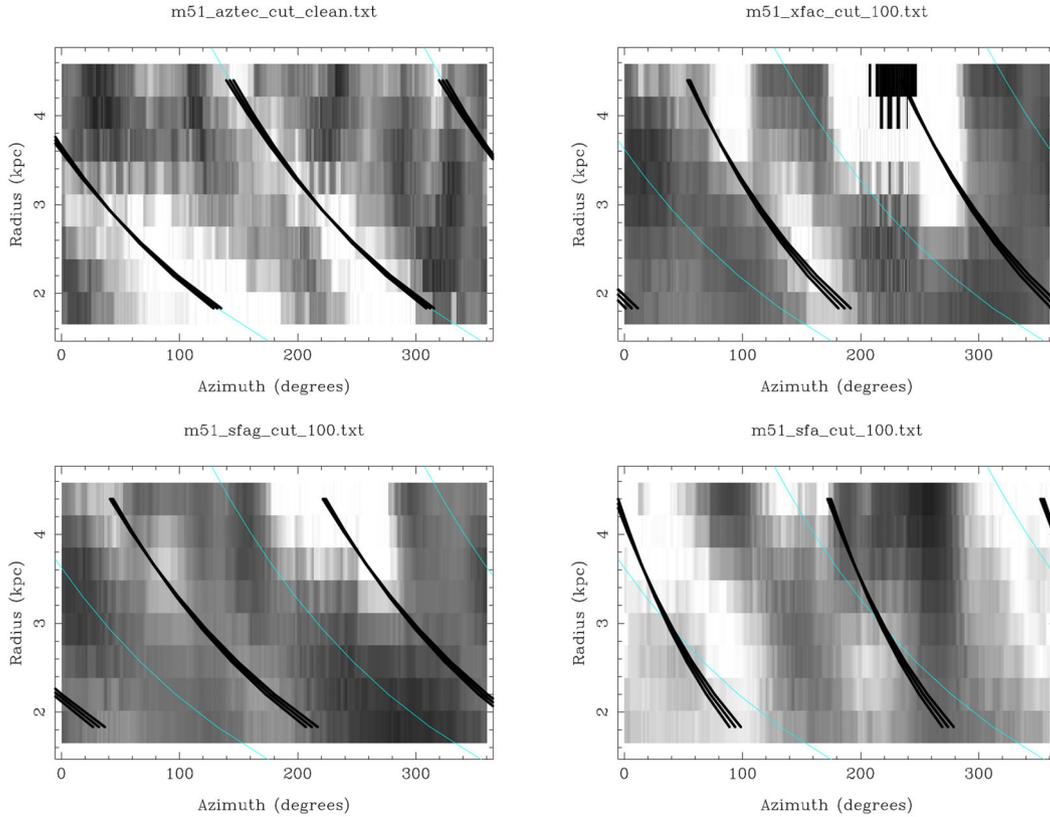


Figure 21. Fourier transform results overplotted on the radius versus theta images for M 51. Upper left: AzTEC image; upper right, X-factor; bottom left, SFAG image; bottom right SFA image. The radius in each panel is in kiloparsecs. Theta is in degrees, where zero is for the west and increases anticlockwise. The minimum to maximum radii in which we conduct the 2D Fourier analysis are 1.8–4.4 kpc. The thin blue lines in the four panels represent the spiral structure we detect in the AzTEC image (pitch angle of 18°). The thick black lines represent the bisymmetrical structure we detect on all images. As clearly seen, the spirals of the X-factor and SFAG images (both with pitch angle of 17°) are rotated with respect to the AzTEC image. This means that the X-factor and SFAG images have higher values in the interarm regions. The spiral structure we detect on the SFA image has a smaller pitch angle (around 13°), but it is in phase with the AzTEC arms.

due to optically thick H I 21-cm emission. Even if that were the case for M 51 and M 83, it would not account for lower limits to X_F that are one or two orders of magnitude higher than the average inner disc value.

In short, none of the alternatives mentioned above are likely to entirely rule out CO-dark gas. Nevertheless, these alternatives themselves are not entirely ruled out either and could partly account for some of the high X-factor values inferred. In any event, more and deeper mapping of CO and other gas tracers of the outer discs of these galaxies is essential for understanding the nature of the dust and gas at these large galactocentric radii.

The existence of CO-dark gas in our Galaxy has been known for a while (see Reach et al. 1998, and references therein) and has been confirmed recently by the Planck Collaboration XIX (2011b). They find $X_F = 2.5 \pm 0.1 X_{20}$ for our Galaxy. Also, as mentioned previously there is much additional evidence for such gas in the Galaxy and external galaxies (see, e.g. Israel 1997a,b; Roman-Duval et al. 2010; Clark et al. 2012; Saintonge et al. 2012; Baes et al. 2014; Langer et al. 2014; Pineda et al. 2014; Smith et al. 2014). So it is quite likely that there is CO-dark gas in both M 51 and M 83. Indeed, the current observations suggest that such gas is present in the outer discs of both M 51 and M 83. The CO-dark gas could account for as much as half of the total gas mass, although the alternative interpretations given above could reduce that fraction by roughly 50 per cent.

5.2 The relationship between the SFR and the gas surface density

In Section 4.4, we examined how the SFR varied with the gas surface density in a number of ways. Comparing between the continuum and spectral line tracers given in Fig. 16, all show higher slopes for the continuum tracers than for the spectral-line tracers. Figs 9 and 11 illustrate that the continuum-derived radial profile of the gas surface density is flatter than that for the spectral-line derived gas surface for large galactocentric radii. Having a smaller range of gas surface densities will naturally increase the slope in the SFR versus gas surface density plots.

Even though Bigiel et al. (2008) used far-UV and $24 \mu\text{m}$ data to estimate the SFR, whereas we used H α and $24 \mu\text{m}$, that is unlikely to account for the difference between their fitted slopes and ours; their Fig. 9 makes such a comparison and the difference in fitted slopes between the two SFR tracers is 10 per cent or less. So if differences in SFR tracers cannot explain the difference between the slope obtained by Bigiel et al. (2008) and those obtained in the current work, then it must be the difference in gas surface density tracer. The continuum tracer used here gives slopes of ~ 2.5 , which is consistent with the eyeball inspection of the far upper-right panel of fig. 8 of Foyle et al. (2012) for the inner ~ 6 kpc radius of M 83.

Estimation of gas surface or column densities is problematic because, as exemplified in the current work, different tracers can yield different results. There are systematics that affect the column

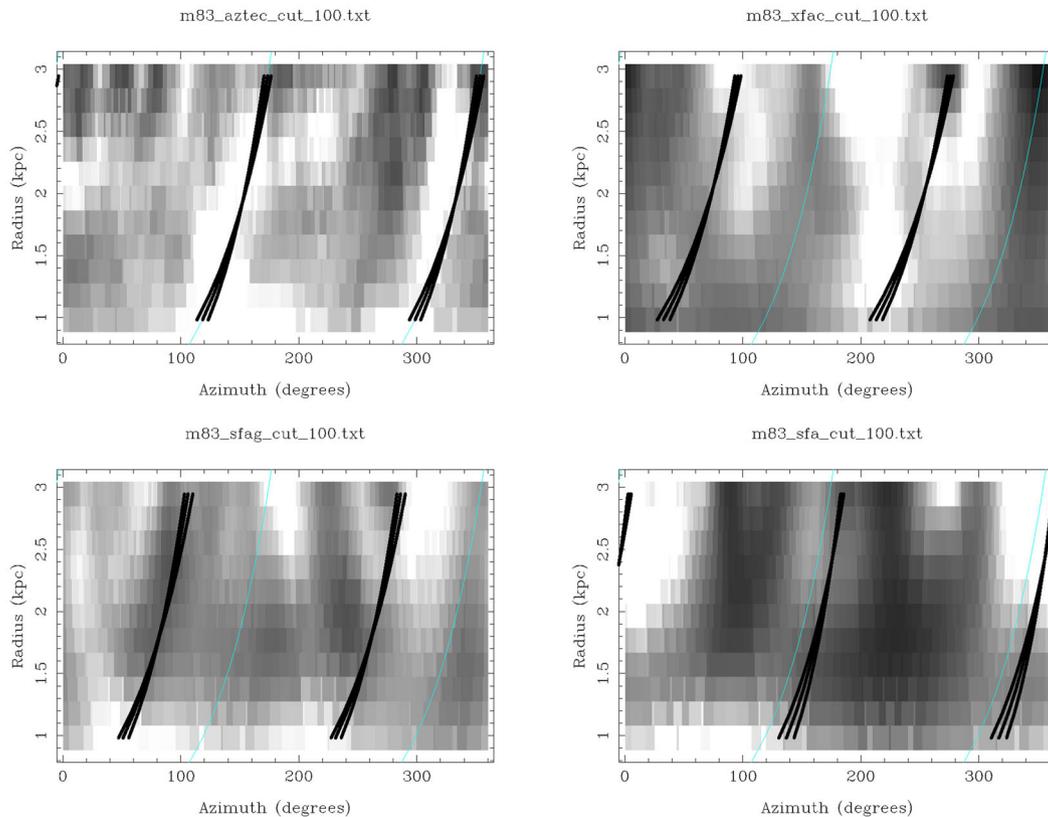


Figure 22. Fourier transform results overplotted on the radius versus theta images for M83. Upper left: AzTEC image; upper right, X-factor; bottom left, SFAG image; bottom right SFA image. The radius in each panel is in kiloparsecs. Theta is in degrees, where zero is for the west and increases anticlockwise. The minimum to maximum radii in which the 2D Fourier analysis is carried out are 1.0 to 3.0 kpc. The thin blue lines in the four panels represent the spiral structure we detect in the AzTEC image (pitch angle of 48°). The thick black lines represent the bisymmetrical structure we detect on all images. The spirals of the X-factor has a pitch angle of 31° , which is similar to that of 36° pitch angle of the SFAG image. The pitch angle of the SFA image is 53° . As found for M51, the X-factor and SFAG spirals are rotated with respect to the AzTEC image, again yielding higher values in the interarm regions.

Table 5. Spiral arm-to-interarm ratios in the inner discs^a of M51 and M83.

Image	M51	M83
Visible ^b	1.3	1.5
1.1 mm	1.7	1.4
X-factor	0.4	0.7
SFA ^c	1.3	2.0
SFAG ^d	0.4	1.0

^aFor galactocentric radii of 1.8-5.5 kpc and 1.0-2.9 kpc for M51 and M83, respectively.

^bIn red light for M51 and *I* band for M83.

^cThe star formation surface density map normalized to the molecular gas surface density determined from the continuum and the H I line.

^dThe star formation surface density map normalized to the molecular gas surface density determined from the CO $J = 1 \rightarrow 0$ line and the X-factor.

density estimates using CO $J = 1 \rightarrow 0$ and different systematics when using infrared and millimetre continuum. These systematics yield the differing slopes and also the observed scatter, which is sometimes quite large with reduced chi-square as high as ~ 40 – 70 . This scatter is partly intrinsic, because the SFR depends on more physical conditions than on just column density. But systematic errors in estimating surface densities also play a role. Corrections

applied systematically to those errors would represent a smooth gradient with galactocentric radius and not simply corrections to a few individual points. For example, observations have inferred X-factor values that are factors of about 5 or more lower than the standard value (e.g. Rickard & Blitz 1985; Israel 1988; Wall et al. 1993; Regan 2000; Paglione et al. 2001) in the centres of external galaxies as well as in the central region of our own Galaxy (Sodroski et al. 1995; Dahmen et al. 1997, 1998). But these centres only represent the central few hundred parsec radii and would thus represent very few points in the plots of Fig. 16. Only a large, smooth gradient in parameters like the X-factor or $\tau_{\text{vd}}/N(\text{H})$ would change the slope uniformly in those plots, and there is no evidence for such.

Nevertheless, there is evidence of a weak dependence of the X-factor on the SFR. Clark & Glover (2015) explored this possibility with their physical models of molecular clouds, finding that $X_F \propto \Sigma_{\text{SFR}}^\gamma$ with $\gamma \simeq 0.5$. The systematically different power-law indices for the dust tracers from those for the gas tracers are reconcilable by adopting that relationship between X_F and Σ_{SFR} . The computation of Appendix F yields $\gamma = -0.38$ and -0.25 for M51 and M83, respectively. These have uncertainties of 3–4 per cent and are thus significantly different from that of the models of Clark & Glover (2015). These are more consistent with the models of Narayanan et al. (2012) (see Clark & Glover 2015). That the M51 and M83 γ values are significantly different from each other argues that there is no universal relationship between X_F and Σ_{SFR} . This is

not surprising given that both of those quantities have complex dependences on the physical structure within the molecular gas.

The value of the slope of the surface densities of SFR versus those of the molecular gas is an indication of the large-scale evolution of a galaxy. As briefly alluded to previously, following the ‘stream’ of points in each of the panels of Fig. 16 from low to high surface densities is equivalent to travelling from large to small galactocentric radii. The $\Sigma_{\text{SFR}}/\Sigma_{\text{H}_2}$ (or alternatively $\Sigma_{\text{SFR}}/[\Sigma_{\text{d}}(\text{gas}) - \Sigma_{\text{H}_1}]$) ratio is really the inverse depletion time of the molecular gas. For slope = 1, $\Sigma_{\text{SFR}}/\Sigma_{\text{H}_2} = \text{constant}$ and the gas is depleted uniformly throughout a galaxy, as pointed out by Bigiel et al. (2008) and found by them and Leroy et al. (2013). However, all the panels of Fig. 16 have slopes >1 , strongly supporting an inside-out depletion of molecular gas in both M 51 and M 83.

The panels of Figs 17 and 18 yield molecular gas depletion times from their centres out to radii of about 8 kpc for M 51 and 6 kpc for M 83. For M 51, the continuum tracers give molecular gas depletion times of about 1.2 Gyr in the centre to 20 Gyr in the outer disc and, with the CO line tracer, these times are 0.8–2.5 Gyr. For M 83, the continuum tracers suggest depletion times of about 0.7 Gyr in the centre to 10 Gyr in the outer disc and, again for CO, these times are about 0.4–5 Gyr. This inside-out evolution of the star formation in the discs of galaxies is supported by the visible-light observations of González Delgado et al. (2014).

At radii from about 8–11 kpc for M 51, the molecular gas depletion times at these radii extend by nearly two orders of magnitude higher than the 20 Gyr estimate for galactocentric radius of 8 kpc. For M 83, this extension of the gas depletion time occurs for radii slightly beyond 6 kpc and is by about 1.5 orders of magnitude.

Figs 17 and 18 are equivalent to fig. 5 of Foyle et al. (2010), displaying the inverse depletion times of the molecular gas in the form of images of entire galaxies. Those images in the current paper (and, to some extent, those of Foyle et al. 2010) are suggestive of spiral structure. Such spiral structure implies an enhancement of the SFR due to the spiral arms beyond that of arms simply collecting and compressing gas and dust. Accordingly, we subjected these images to spiral arm Fourier analysis and found only a weak spiral structure with arm/interarm ratios usually within factors of 2 of unity. The right-hand panels were found to have a spiral structure with high interarm values of $\Sigma_{\text{SFR}}/\Sigma_{\text{H}_2}$ and lower arm values, which is corrected when accounting for the higher interarm X_{F} (see Section 4.3). In contrast, the left-hand panels were found to have a spiral structure with high $\Sigma_{\text{SFR}}/[\Sigma_{\text{d}}(\text{gas}) - \Sigma_{\text{H}_1}]$ on the arms and lower values between the arms; the arm/interarm ratios are 1.3 for M 51 and 2.0 for M 83. These arm/interarm ratios are at least partly explained by the higher $\tau_{\text{vd}}/N(\text{H})$ between the arms than in the arms (see Appendix B1) – 18 per cent higher for M 51 and 8 per cent higher for M 83.

So, conservatively speaking, even the continuum tracer in our work confirms the work of Foyle et al. (2010) that spiral arms only enhance the star formation rate because of increasing the surface density of gas and dust with no additional enhancement. The uncertainties in the current work do not permit completely ruling out such an enhancement, but suggest that any such enhancement would be small (i.e. a factor of $\lesssim 2$). That enhancement, should it be real, could be accounted for by orbit-crowding in the spiral arms raising the inverse depletion time of the gas in those arms (e.g. Moore et al. 2012).

One key question is whether the conclusions are still valid if the diffuse emission is removed (Foyle et al. 2013). But the lower spatial resolution of the gas tracer observations impede determination of the surface densities associated only with the star-forming

regions. So a surface-density versus surface-density plot is difficult to create (see Foyle et al. 2013), making comparison in the context of much previous work difficult. Removing the diffuse emission might still result in the slope of the $\log(\text{SFR surface density})$ versus $\log(\text{gas surface density})$ still being higher when a dust-continuum tracer is used in place of a gas-line tracer (see Fig. 16), due to the CO-dark gas not traced by spectral lines or maybe due to X_{F} varying with the SFR. Also, the inside-out galactic-scale evolution of star formation, as indicated by the superlinear slopes, is likely still valid due to the support of independent work (González Delgado et al. 2014).

Shetty et al. (2013) use hierarchical Bayesian linear regression on the observational data, finding that no one S-K relation holds for all galaxies. This is consistent with the current work where no one power law applies to either M 51 or M 83. Indeed, any simple power-law fit is inapplicable given the poor quality of fits in Fig. 16. Not only are there different slopes in different galaxies, but there are also different offsets, particularly in post-starburst galaxies (see French et al. 2015). And, if no one such relation holds for all galaxies, then the S-K relation does not represent a universal physical law of star formation (see Lada et al. 2013). Perhaps the S-K relation for a disc galaxy is a measure of how the inverse depletion time of the molecular gas varies radially in that disc galaxy (e.g. inside-out or outside-in star formation), which, in turn, is affected by the many properties of, and processes in, the disc of that particular galaxy (see, for example Hughes 2013; Meidt et al. 2013; Schinnerer et al. 2013; Colombo et al. 2014).

6 SUMMARY AND CONCLUSIONS

The spiral galaxies M 51 and M 83 were observed with the bolometer array AzTEC on the JCMT in the 1.1 mm continuum at 20 arcsec spatial resolution. The extended emission, including the interarm emission and exponential discs, was faithfully recovered in the final maps out to galactocentric radii of more than 12 kpc for both galaxies. This was verified by simulations that show that only small corrections are necessary for the M 51 image and negligible corrections for that of M 83. The 1.1 mm-continuum fluxes are 5.6 ± 0.7 and 9.9 ± 1.4 Jy for M 51 and M 83, respectively. The uncertainties are largely due to that of the calibration.

These images were combined with the 160 μm image of *Spitzer* to obtain dust temperatures and column densities. This approach was adopted, rather than using dust models, for two reasons. One reason was to have an independent test of the models (see Groves et al. 2015). The other reason is that multiwavelength far-IR data at the long wavelengths (i.e. $\lambda \gtrsim 100 \mu\text{m}$) at which the bulk of the dust mass radiates is not available for the outer discs of M 51 and M 83, other than the 160 μm data of *Spitzer*.

Another model-independent approach was to constrain, at least roughly, the $\tau_{\text{vd}}/N(\text{H})$ at 1.1 mm by the observations, rather than simply adopting a dust mass absorption coefficient, $\kappa_{\text{v}}(1.1 \text{ mm})$. Gas column densities were estimated by calibrating against H I-dominant positions to estimate the dust optical depth to gas column density ratio $\tau_{\text{vd}}/N(\text{H})$. The method of calibrating against the H I-dominant positions was improved by crudely estimating the effects of the CO-dark gas (see below). Neither galaxy has a strong radial variation in the gas surface density beyond galactocentric radii of about 3 kpc. Out to a galactocentric radius of 14 kpc, the best estimate of the mass of gas in M 51 is $9.4 \times 10^9 M_{\odot}$. Out to 12 kpc in M 83, this best estimate is $7.2 \times 10^9 M_{\odot}$. (See adopted distances in Table 1.)

Pre-existing maps of CO $J = 1 \rightarrow 0$ permitted the creation of maps of the $N(\text{H}_2)/I(\text{CO})$ or X-factor for both M 51 and M 83

out to galactocentric radii of 6–8 kpc. Both galaxies have X-factor values that are higher in the interarm than in the arms by a factor of ~ 1.5 –2. In the central few kiloparsecs of M51, this interarm/arm ratio rises to ~ 3 . In M51, there is no significant radial variation of X_F . In M83, however, there is evidence at the many- σ level of radial variation of the X-factor by factors of 2 to 3, where the central 2 kpc radius has a roughly flat X-factor which declines to a minimum at about 4 kpc. Within galactocentric radii ≤ 6 –8 kpc, the spatially averaged X-factor is about $1 X_{20}$.

Beyond the outer radius of the X-factor map for each galaxy, comparison of the radial profile of the gas surface density derived from the continuum with those of the surface brightnesses of CO and H₁ permits estimates of lower limits of the X-factor that reach one to two orders of magnitude higher than the inner disc values. This suggests the existence of CO-dark molecular gas in the outer discs of M51 and M83, although alternative explanations are only partly ruled out. Nevertheless, these alternatives do not entirely account for the high X-factor values in the outer disc.

A 2D Fourier analysis of the spiral structure at 1.1 mm and at visible (or near visible) wavelengths revealed that the spiral structure in red light and that in the 1.1 mm continuum and were the same in M51. For M83, the spiral structure in *I* band compared with that in the 1.1 mm continuum showed that the bar’s effect in M83 is conspicuous in *I* band and not at 1.1 mm. These results suggest that the spiral density wave in M51 is influencing the interstellar medium and stars similarly, while the bar potential in M83 has a different influence on the interstellar medium from that on the stars.

Log–log plots of the SFR surface densities against those of the gas traced by spectral lines (i.e. of H₁ and CO) had slopes of ~ 1.5 whether total gas or just molecular gas surface density. For the plots with gas surface densities traced by the continuum emission, the slopes were ~ 2.5 whether total gas (using continuum only) or just molecular (using continuum with H₁ subtracted) gas surface densities. These plots, especially with the continuum tracers, show a threshold gas surface density at which the SFR rises by two or more orders of magnitude. The existence of this threshold gas surface density is insensitive to within a factor of ~ 3 for the adopted $\tau_{\text{vd}}/N(\text{H})$. The value of this threshold density is $\sim 15 M_{\odot} \text{pc}^{-2}$. This threshold is somewhat less conspicuous in the spectral line tracers than for the continuum tracers.

The fitted slopes suggest that the depletion of the molecular gas occurs first at small galactocentric radii and then at increasing radii in an inside-out galactic evolution. This is seen more clearly in maps of the ratio of the surface densities of the SFR to that of the molecular gas. For both these galaxies and the continuum tracer, the molecular gas depletion time in the centres is about 1 Gyr, rising at radii of 6–8 kpc to around 10–20 Gyr. Further out, the depletion times rise by one or two orders of magnitude. The spectral line tracer, i.e. CO $J = 1 \rightarrow 0$, suggests molecular gas depletion times in the outer discs that are appreciably less than 10–20 Gyr.

The images of the inverse depletion time show signs of spiral structure. Superficially, this suggests that spiral arms affect the SFR beyond just heightening the gas surface density. However, correcting for the X-factor spatial variation and for the spatial variation of the $\tau_{\text{vd}}/N(\text{H})$ removes or nearly removes such spiral structure. This apparently confirms the result of (Foyle et al. 2010) that the arms merely heighten the SFR in the same proportion as they heighten the gas surface density. Greater spatial resolution is required for confirming this result.

In the future, we need deeper mapping of molecular tracers in the outer discs of these spiral galaxies. As well, a better method of calibrating the $\tau_{\text{vd}}/N(\text{H})$ at millimetre wavelengths is needed. So far, the method employed in the current work is functional, but only crudely. Either a new method or refinement of the method described here is necessary.

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REFERENCES

- Aniano G. et al., 2012, *ApJ*, 756, 138
 Baes M. et al., 2014, *MNRAS*, 444, L90
 Baillard A. et al., 2011, *A&A*, 532, 74
 Bendo G. J. et al., 2012, *MNRAS*, 419, 1833
 Bigiel F., Leroy A., Walter F., Brinks E., de Blok W. J. G., Madore B., Thornley M. D., 2008, *AJ*, 136, 2846
 Blasco-Herrera J. et al., 2010, *MNRAS*, 407, 2519
 Block D. L., Puerari I., 1999, *A&A* 342, 627
 Boselli N. et al., 2010, *A&A*, 518, L61
 Boselli N. et al., 2012, *A&A* 540, 54
 Bourne N. et al., 2012, *MNRAS*, 421, 3027
 Bresolin F., Garnett D. R., Kennicutt R. C., 2004, *ApJ*, 615, 228
 Bresolin F., Ryan-Weber E., Kennicutt R. C., Goddard Q., 2009, *ApJ*, 695, 580
 Calzetti D. et al., 2007, *ApJ*, 666, 870
 Calzetti D., Liu G., Koda J., 2012, *ApJ*, 752, 98
 Clark P. C., Glover S. C. O., 2014, *MNRAS*, 444, 2396
 Clark P. C., Glover S. C. O., 2015, *MNRAS*, 452, 2057
 Clark P. C., Glover S. C. O., Klessen R. S., Bonnell I. A., 2012, *MNRAS*, 424, 2599
 Colombo D. et al., 2014, *ApJ*, 784, 4
 Considère S., Athanassoula E., 1988, *A&AS* 76, 365
 Cooper E. M. et al., 2012, *ApJ*, 755, 165
 Cortese L. et al., 2012, *A&A* 540, 52
 Cortese L. et al., 2014, *MNRAS*, 440, 942
 Crosthwaite L. P., Turner J. L., Buchholz L., Ho P. T. P., Martin R. N., 2002, *AJ*, 123, 1892
 Dahmen G. et al., 1997, *A&AS*, 126, 197
 Dahmen G., Hüttemeister S., Wilson T. L., Mauersberger R., 1998, *A&A* 331, 959.
 Dale D. A. et al., 2009, *ApJ*, 703, 517
 Dettmar R.-J., Heithausen A., 1989, *ApJ*, 344, L61
 Dobbs C. L., 2015, *MNRAS*, 447, 3390
 Draine B. T., 2003, *ARA&A*, 41, 241
 Dunne L. et al., 2011, *MNRAS*, 417, 1510
 Eales S. et al., 2010, *A&A*, 518, L62
 Evans N. J., 1980, in Andrew B. H., ed., *Proc. IAU Symp. 87, Interstellar Molecules*. Reidel, Dordrecht, p. 1
 Evans N. J., 1999, *ARA&A*, 37, 311
 Foyle K., Rix H.-W., Walter F., Leroy A. K., 2010, *ApJ*, 725, 534
 Foyle K. et al., 2012, *MNRAS*, 421, 2917
 Foyle K. et al., 2012, *MNRAS*, 432, 2182
 French K. D., Yang Y., Zabludoff A., Narayanan D., Shirley Y., Walter F., Smith J.-D., Tremonti C. A., 2015, *ApJ*, 801, 1
 Fritz J. et al., 2012, *A&A*, 546, 34
 Fukui Y. et al., 1999, *PASJ*, 51, 745

- González Delgado R. M. et al., 2014, *A&A*, 562, 47
 Gordon K. D. et al., 2010 *A&A* 518, L89
 Groves B. A. et al., 2015, *ApJ*, 799, 96
 Hagiwara Y., Henkel C., Menten K. M., Nakai N., 2001, *ApJ*, 560, 37
 Heiderman A., Evans N. J., Allen L. E., Huard T., Heyer M., 2010, *ApJ*, 723, 1019
 Helfer T. T., Thornley M. D., Regan M. W., Wong T., Sheth K., Vogel S. N., Blitz L., Bock D. C.-J., 2003, *ApJS*, 145, 259
 Hughes A., 2013, *ApJ*, 779, 46
 Israel F. P., 1988, in *Wolstencroft R. D., Burton W. B., eds, Millimetre and Submillimetre Astronomy, Summer School, Stirling, Scotland. Kluwer, Dordrecht*, p. 281
 Israel F. P., 1997a, *A&A*, 317, 65
 Israel F. P., 1997b, *A&A*, 328, 471
 James A., Dunne L., Eales S., Edmunds M., 2002, *MNRAS*, 335, 753
 Kalnajs A., 1975, in *Weliachew L., ed., La Dynamique des Galaxies Spirales*. p. 103
 Karachentsev I. D. et al., 2002, *A&A*, 385, 21
 Kennicutt R. C., 1989, *ApJ*, 344, 685
 Kennicutt R. C., 1998, *ApJ*, 498, 541
 Kennicutt R. C. et al., 2003, *PASP*, 115, 928
 Kirkpatrick A. et al., 2014, *ApJ*, 789, 130
 Komugi S. et al., 2011, *PASJ*, 63, 1139
 Kutner M. L., 1984, *Fundam. Cosm. Phys.*, 9, 233
 Lada C. J., Lombardi M., Roman-Zuniga C., Forbrich J., Alves J. F., 2013, *ApJ*, 778, 133
 Lagos C. d. P., Bayet E., Baugh C. M., Lacey C. G., Bell T. A., Fanidakis N., Geach C. G., 2012, *MNRAS*, 426, 2142
 Langer W. D., Velusamy T., Pineda J. L., Willacy K., Goldsmith P. F., 2014, *A&A*, 561, 122
 Leroy A. K. et al., 2013, *ApJ*, 146, 19
 Li A., Draine B. T., 2001, *ApJ*, 554, 778
 Lombardi M., Bouy H., Alves J., Lada C. J., 2014, *A&A*, 566, 45
 Lundgren A. A., Wiklund T., Olofsson H., Rydbeck G., 2004, *A&A*, 413, 505
 Madden S. C., Poglitsch A., Geis N., Stacey G. J., Townes C. H., 1997, *ApJ*, 483, 200
 Magnelli B. et al., 2012, *A&A*, 548, 22
 Maloney P. R., Black J. H., 1988, *ApJ*, 325, 389
 Meidt S. E. et al., 2013, *ApJ*, 779, 45
 Meijerink R., Tilanus R. P. J., Dullemond C. P., Israel F. P., van der Werf P. P., 2005, *A&A*, 430, 427
 Meurer G. R. et al., 2006, *ApJS*, 165, 307
 Miller E. D., Bregman J. N., Wakker B. P., 2009, *ApJ*, 692, 470
 Moore T. J. T., Urquhart J. S., Morgan L. K., Thompson M. A., 2012, *MNRAS*, 426, 701
 Nakai N., Kuno N., 1995, *PASJ*, 47, 761
 Narayanan D., Krumholz M., Ostriker E. C., Hernquist L., 2011, *MNRAS*, 418, 664
 Narayanan D., Krumholz M., Ostriker E. C., Hernquist L., 2012, *MNRAS*, 421, 3127
 Neininger N., Beck R., Sukumar S., Allen R. J., 1993, *A&A*, 274, 687
 Nishiyama K., Nakai N., Kuno N., 2001, *PASJ*, 53, 757
 Ossenkopf V., Henning T., 1994, *A&A*, 291, 943
 Paglione T. A. D. et al., 2001, *ApJs*, 135, 183
 Pan H.-A., Kuno N., Hirota A., 2014, *PASJ*, 66, 27
 Papadopolis P. P., van der Werf P., Xilouris E., Isaak K. G., Gao Y., 2012, *ApJ*, 751, 10
 Parkin T. J. et al., 2013, *ApJ*, 776, 65
 Pineda J. L., Langer W. D., Goldsmith P. F., 2014, *A&A*, 570, 121
 Planck Collaboration, I 2011a, *A&A*, 536, 16
 Planck Collaboration XIX, 2011b, *A&A*, 536, 19
 Planck Collaboration XXV, 2011c, *A&A*, 536, 25
 Puerari I., Dottori H. A., 1992, *A&AS*, 93, 469
 Reach W. T., Wall W. F., Odegard N., 1998, *ApJ*, 507, 507
 Regan M. W., 2000, *ApJ*, 541, 142
 Regan M. W., Thornley M. D., Helfer T. T., Sheth K., Wong T., Vogel S. N., Blitz L., Bock D. C.-J., 2001, *ApJ*, 561, 218
 Regan M. W. et al., 2006, *ApJ*, 652, 1112
 Rickard L. J., Blitz L., 1985, *ApJ*, 292, L57
 Roman-Duval J. et al., 2010, *A&A*, 518 L74.
 Rots A. H., Bosma A., van der Hulst J. M., Athanassoula E., Crane P. C., 1990, *AJ*, 100, 387
 Rowlands K. et al., 2012, *MNRAS*, 419, 2545
 Saintonge A. et al., 2012, *ApJ*, 758, 73
 Sandstrom K. M. et al., 2013, *ApJ*, 777, 5
 Santini P. et al., 2014, *A&A*, 562, 30
 Schinnerer E. et al., 2013, *ApJ*, 779, 42
 Schmidt M., 1959, *ApJ*, 129, 243
 Shetty R., Vogel S. N., Ostriker E. C., Teuben P. J., 2007, *ApJ*, 665, 1138
 Shetty R., Glover S. C., Dullemond C. P., Klessen R. S., 2011a, *MNRAS*, 412, 1686
 Shetty R., Glover S. C., Dullemond C. P., Ostriker E. C., Harris A. I., Klessen R. S., 2011b, *MNRAS*, 415, 3253
 Shetty R., Kelly B. C., Bigiel F., 2013, *MNRAS*, 430, 288
 Smith M. W. L. et al., 2012, *ApJ*, 748, 123
 Smith R. J., Glover S. O. C., Clark P. C., Klessen R. S., Springel V., 2014, *MNRAS*, 441, 1628
 Sodroski T. J. et al., 1995, *ApJ*, 452, 262
 Talbot R. J., Jr, Jensen E. B., Dufour R. J., 1979, *ApJ*, 229, 91
 Tilanus R. P. J., Allen R. J., 1993, *A&A*, 274, 707
 Tully R. B., 1974, *ApJS*, 27, 449
 Wall W. F., 2007, *MNRAS*, 379, 674
 Wall W. F., Jaffe D. T., Bash F. N., Israel F. P., Maloney P. R., Baas F., 1993, *ApJ*, 414, 98
 Walter F., Brinks E., de Blok W. J. G., Bigiel F., Kennicutt R. C., Thornley M. D., Leroy A., 2008, *AJ*, 2563, 136
 Wilson G. et al., 2008, *MNRAS*, 386, 807

APPENDIX A: FLUX DETERMINATION

The outer edge to the 1.1 mm continuum emission of M 51 lies at a galactocentric radius of about 14 kpc and for M 83 it is about 12 kpc. The derived fluxes are 5.6 ± 0.7 Jy for M 51 and 9.9 ± 1.4 Jy for M 83 to those outer edges. The uncertainties include the calibration uncertainty of 13 per cent and the uncertainties due to the somewhat arbitrary choices for outer boundaries. The simulations were used to check for systematics in the flux determination, thereby providing correction factors for the derived fluxes. For M 51, this correction factor was 1.5 for M 51 and 0.9 for M 83. These corrections were included in the previously quoted flux values.

APPENDIX B: DETERMINATION OF SURFACE DENSITIES AND MASSES

The 1.1 mm continuum surface-brightness maps of Figs 1 and 2 were ratioed with the *Spitzer*/MIPS 160 μm , effectively 155.9 μm , maps of M 51 and M 83 (Dale et al. 2009) to produce temperature maps under the assumption that the dust emissivity index, β , was 2.0. As stated in the Introduction, this was the most appropriate value for β as found from the observations of M 83 by Foyle et al. (2012). In practice, the adopted value for β will have only a small effect on the derived gas column densities, provided that those derived column densities are calibrated against observed gas column densities. By using those positions where atomic gas dominated the column densities to calibrate the $\tau_{\text{vd}}/N(\text{H})$ for the continuum-derived column densities, it was found that the column densities derived from the continuum data for $\beta = 1.5$ –2.5 did not deviate by more than 30 per cent from those for $\beta = 2.0$. Consequently, observations at only two wavelengths are necessary for a reasonable approximation of the gas surface density map.

The $\tau_{\text{vd}}/\text{N}(\text{H})$ derived from the H I comparisons are $2.2\text{--}2.3 \times 10^{-26} \text{ cm}^2$ for M 51 and M 83, respectively. These values lead to gas mass estimates that appear to be too low (see Sections 4.3, 5.1, and below in this appendix for details). Foyle et al. (2012) simply adopted the dust mass-absorption coefficient of Li & Draine (2001) and Draine (2003), which is a factor of ~ 6 smaller than the values determined here. Foyle et al. (2012) also determined a higher dust-to-gas mass ratio than that adopted here. This is equivalent to adopting a $\tau_{\text{vd}}/\text{N}(\text{H})$ that is factors of 5.0–5.3 smaller than that derived from the H I comparisons. Because we adopt a constant dust-to-gas ratio, the two $\tau_{\text{vd}}/\text{N}(\text{H})$ cases are referred to as the high- κ_{ν} (from H I comparisons) and low- κ_{ν} (Li & Draine 2001; Draine 2003) cases. Given the low masses that result from the high- κ_{ν} case, the H I comparisons must be appropriately modified. The H I comparisons were repeated after removing upper outliers that are apparently unreliable due to possible undetected gas. This results in an intermediate case, where a $\tau_{\text{vd}}/\text{N}(\text{H})$ that is a factor of 2 lower than that from the simple atomic gas calibration method is found. Unless otherwise stated, all masses and surface densities quoted are for this intermediate- $\tau_{\text{vd}}/\text{N}(\text{H})$ (most realistic) case. This approach yields a roughly 50 per cent uncertainty in column density and mass estimates, using $\tau_{\text{vd}}/\text{N}(\text{H}) = 1.1 \times 10^{-26} \text{ cm}^2$.

The surface brightness at frequency ν , $I_{\nu}(T_{\text{d}})$, is related to the dust-derived gas column density, $\text{N}_{\text{d}}(\text{H})$, by

$$I_{\nu} = \mu m_{\text{H}} x_{\text{d}} \kappa_{\nu} \text{N}_{\text{d}}(\text{H}) B_{\nu}(T_{\text{d}}), \quad (\text{B1})$$

where μ is the mean atomic weight per hydrogen atom, m_{H} is the mass of a hydrogen atom, x_{d} is the dust-to-gas mass ratio, and κ_{ν} is the dust mass absorption coefficient. The mass-absorption coefficient varies with frequency as follows

$$\kappa_{\nu} = \kappa_{\nu_0} \left(\frac{\nu}{\nu_0} \right)^{\beta}, \quad (\text{B2})$$

in which κ_{ν_0} is dust mass-absorption coefficient at a reference frequency ν_0 . The dust temperature, T_{d} , is of course estimated by evaluating expression (B1) at frequency ν_1 , corresponding to the wavelength of 1.1 mm for the JCMT/AzTEC observations, and again at frequency ν_2 , corresponding to the wavelength of 155.9 μm for the *Spitzer*/MIPS observations, and taking the ratio of the two.

As well as using the maps at the two wavelengths (i.e. the JCMT/AzTEC at 1.1 mm and the *Spitzer*/MIPS map at 155.9 μm), there are many additional details involved in estimating the column densities, $\text{N}_{\text{d}}(\text{H})$. These include estimating the value of $\tau_{\text{vd}}/\text{N}(\text{H})$ (or the product of x_{d} and κ_{ν}), convolving the AzTEC 1.1 mm map to the resolution of the *Spitzer*/MIPS 155.9 μm map, estimating the noise levels in the maps, characterizing any systematic effects introduced into the T_{d} and $\text{N}_{\text{d}}(\text{H})$ maps due to the observations and processing of the AzTEC data, and any relevant colour corrections to the continuum data. In particular, a constant offset correction of $5.5 \times 10^{20} \text{ H} \text{---} \text{nuclei cm}^{-2}$ must be added to the column densities for M 51 as dictated by the simulations to correct for inadequately recovering the large-scale emission (but was not applied to the figures).

B1 The reliability of gas surface densities and masses

There are a couple issues that should be addressed in discussing the reliability of gas surface densities and masses as inferred from dust emission in the current work. One issue is to what level calibration differences between that of *Spitzer*/MIPS and that of *Herschel*/PACS for their 160 μm data would affect the results and conclusions of the current work (see Aniano et al. 2012). Another issue

is the validity of the simple approach used here instead of using a more detailed dust model (e.g. Li & Draine 2001). As for the former issue, Aniano et al. (2012) compared the 160 μm data for both instruments and characterized their systematic differences. Depending on the calibration scheme used, the *Herschel*/PACS 160 μm intensities were either 1.6 per cent or 25 per cent systematically higher than those for *Spitzer*/MIPS. Scaling up the 160 μm intensities used in the current work by 25 per cent would increase the computed dust temperatures, but by differing amounts depending on initial computed temperature. Most of the those temperatures were determined to be between 10 and 23 K. After the hypothetical correction, those would be 10.3 and 25 K. If there were no calibration of the column densities against other data, those corrected temperatures would require the column densities to be corrected downwards by 5 per cent and 11 per cent, respectively. The warmer positions in both galaxies tend to have higher column densities. So the higher column densities would be corrected downwards by more than the lower column densities, on average. This reduces the dynamic range of the determined column densities by roughly 6 per cent. The overall scaling of those column densities would remain unchanged because of the calibration against H I column densities. So, such a correction, were it necessary, would not appreciably change the current results.

As for the second issue, the models of Li & Draine (2001) combined with data at mid-IR and far-IR wavelengths can provide estimates of a number of parameters. For the current work, however, there are two reasons why such an approach was not used. One reason is that, as mentioned in the introduction, the AzTEC 1.1 mm map has more spatial coverage than the *Herschel* data for both M 51 and M 83. *Spitzer* does have sufficient spatial coverage for comparison with the AzTEC data, but does not have the long-wavelength data that probes the bulk of the dust mass, other than at 160 μm . So, to consistently treat each entire galaxy, the AzTEC 1.1 mm data were combined with the *Spitzer* 160 μm data for both M 51 and M 83. Another reason that the more model-dependent approach was not used is that we need to have model-independent tests that can confirm or refute the validity of the models, as stated recently by Groves et al. (2015). The current work, for example, finds a dust opacity to gas column density ratio $\tau_{\text{vd}}/\text{N}(\text{H})$ at 1.1 mm that is about double that inferred from Li & Draine (2001). But temporarily adopting the $\tau_{\text{vd}}/\text{N}(\text{H})$ of Li & Draine (2001) and adopting the appropriate distance yields nearly identical masses for both the more sophisticated approach (i.e. in Cooper et al. 2012) and the simpler approach (i.e. the current work) for the NGC 5194 field. This strongly suggests that the simpler approach of the current work is reasonable at recovering some basic dust properties.

The dust optical depth to gas column density ratio, $\tau_{\text{vd}}/\text{N}(\text{H})$, and especially the dust mass absorption coefficient, κ_{ν} , can be quite uncertain, ranging from $\kappa_{\nu}(1.1 \text{ mm}) = 0.06$ to $1.4 \text{ cm}^2 \text{ g}^{-1}$. (see the following for examples of disparate values: Ossenkopf & Henning 1994; Li & Draine 2001; James et al. 2002; Draine 2003; Eales et al. 2010; Dunne et al. 2011). Accordingly, attempting to measure $\tau_{\text{vd}}/\text{N}(\text{H})$ observationally could narrow the uncertainty. This is most easily observed in gas dominated by atomic hydrogen because the gas column densities are easy to determine. Another approach was followed by Eales et al. (2010) for observations of the galaxies M 99 and M 100 in which they plotted the SFR surface density, Σ_{SFR} , against gas surface density and against that for dust, adjusting the κ_{ν} until both surface densities agree at large Σ_{SFR} . This resulted in a κ_{ν} value corresponding to the unusually low $\kappa_{\nu}(1.1 \text{ mm}) = 0.06 \text{ cm}^2 \text{ g}^{-1}$ mentioned above.

The large $\tau_{\text{vd}}/N(\text{H})$ initially derived in the current work was from comparing the observed optical depths at 1.1 mm with the observed gas column densities in regions dominated by atomic gas. The Planck Collaboration XIX (2011b) and Planck Collaboration XXV (2011c) observed $\tau_{\text{vd}}/N(\text{H}) = 5.2 \times 10^{-26} \text{ cm}^2$ at 857 GHz and $\tau_{\text{vd}}/N(\text{H}) = 1.1 \times 10^{-25} \text{ cm}^2$ at 250 μm in the dust associated with the H I gas within the Galaxy. The Planck Collaboration XXV (2011c) also observed $\tau_{\text{vd}}/N(\text{H}) = 2.3 \times 10^{-25} \text{ cm}^2$ at 250 μm for the dust associated with molecular gas, which is about double that observed for dust associated with atomic gas, remembering that the $N(\text{H})$ is the column density of hydrogen nuclei. At a wavelength of 1.1 mm, these numbers correspond to $6.6 \times 10^{-27} \text{ cm}^2$ and $7.9 \times 10^{-27} \text{ cm}^2$ for dust associated with H I and $\tau_{\text{vd}}/N(\text{H}) = 1.6 \times 10^{-26} \text{ cm}^2$ for that associated with H_2 (using a spectral emissivity index of $\beta = 1.8$ for the Galaxy, e.g. Planck Collaboration XIX 2011b; Planck Collaboration XXV 2011c). The current work initially finds $\tau_{\text{vd}}/N(\text{H}) = 2.2 \times 10^{-26} \text{ cm}^2$ for M 51 and $2.3 \times 10^{-26} \text{ cm}^2$ for M 83 for dust along lines of sight with gas column densities dominated by H I . This is more than double the values observed in our Galaxy and are even 40 per cent larger than the dust associated with H_2 in our Galaxy. This large $\tau_{\text{vd}}/N(\text{H})$ is suspect and requires a re-evaluation of the calibration of $\tau_{\text{vd}}/N(\text{H})$ at 1.1 mm.

One explanation for the observed large $\tau_{\text{vd}}/N(\text{H})$ would be CO-dark gas. The presence of such gas can be crudely tested. This is done by starting with dust-derived gas column densities, $N_{\text{d}}(\text{H})$, that are only roughly calibrated – i.e. to within a factor of a few. The difference $N_{\text{d}}(\text{H}) - N(\text{H I})$ then provides a rough measure of molecular gas column density. The ratio $(N_{\text{d}}(\text{H}) - N(\text{H I})) / (2 X_{\text{F}} I(\text{CO}))$ (where the X_{F} is the average used for each galaxy) is a crude measure of the amount of dark gas – a dark ratio. Then $\tau_{\text{vd}}/N(\text{H})$ can be recomputed using only those positions with dark ratios less than some upper limit. Choosing a dark ratio upper limit as high as 100 reduces the derived $\tau_{\text{vd}}/N(\text{H})$ by a few per cent, while choosing unity yields a $\tau_{\text{vd}}/N(\text{H})$ that is a factor of about 4 lower. The latter is risky given that this could unnecessarily exclude too many positions with very little dark gas. Accordingly, a more conservative correction factor would be around 2.

Reducing the $\tau_{\text{vd}}/N(\text{H})$ value by this factor scales up the column densities by the same factor and the X-factor scales in proportion to $N_{\text{d}}(\text{H}) - N(\text{H I})$, or up by a factor of ~ 3 for both M 51 and M 83. The best estimates of these values are hence $\tau_{\text{vd}}/N(\text{H}) \simeq 1 \times 10^{-26} \text{ cm}^2$, corresponding to $\kappa_{\nu}(1.1 \text{ mm}) \simeq 0.7 \text{ cm}^2 \text{ g}^{-1}$ for $\mu x_{\text{d}} = 0.01$ for both M 51 and M 83. These are well within a factor of 2 of the values found by the Planck Collaboration XIX (2011b) and Planck Collaboration XXV (2011c) for our Galaxy. The best estimates for the X-factors for M 51 and M 83 are around $1 X_{20}$. Table 3 displays the best estimated values for the gas masses and the X-factors. That table shows that roughly half of the total mass of the gas is unaccounted for using spectral lines. This is discussed further in Section 5.1.

Possible spatial variations in $\tau_{\text{vd}}/N(\text{H})$ should also be considered. Both M 51 and M 83 show decreases in $\tau_{\text{vd}}/N(\text{H})$ by about 15–20 per cent in going from galactocentric radii of 1 or 2 kpc to about 6 or 7 kpc. Correcting for the CO-dark gas changes this result. The problem is that the number of points used in the comparisons becomes very small (i.e. ~ 10) when the upper limit to the dark ratio is reduced. Such small numbers do not give reliable estimates of $\tau_{\text{vd}}/N(\text{H})$. So the radial spatial variation of $\tau_{\text{vd}}/N(\text{H})$ is uncertain, although there is no clear evidence that it has a strong radial variation. In going from interarm positions to spiral arm positions, the $\tau_{\text{vd}}/N(\text{H})$ in M 51 decreases by 18 per cent. In M 83, this decrease

is 8 per cent. Again, correcting for the CO-dark gas changes these results to 25 per cent for M 51 and an 11 per cent increase for M 83, although the number of points available for the estimation is not really sufficient.

APPENDIX C: DETERMINATION OF THE X-FACTOR MAPS

The X-factor, X_{F} , can be straightforwardly computed from

$$X_{\text{F}} = \frac{N_{\text{d}}(\text{H}) - N(\text{H I})}{2 I(\text{CO})}. \quad (\text{C1})$$

Given that $N_{\text{d}}(\text{H})$, $N(\text{H I})$, and $I(\text{CO})$ are all maps, X_{F} is also a map, but not a very extended map because of the mounting uncertainties that result from the numerous computations with pre-existing data required to reach X_{F} . Both the M 51 and M 83 X_{F} maps do not extend beyond about 7 kpc from the centre of each map, mostly because the CO maps do not extend further than that. Figs 12–15 show the X-factor maps and radial profiles for M 51 and M 83. The mean values of the X-factor listed in Table 4, which we call X_{FM} , are the $1/\sigma(X_{\text{F}})^2$ -weighted means of each X_{F} map and only really applies to the central 7 kpc radius in each galaxy.

Given that the CO $J = 1 \rightarrow 0$ emission falls to the noise level at a radius of about 7–8 kpc, we are unable to compute X_{F} beyond that in the outer discs. Nevertheless, we can use the radial curves of Figs 9 and 11 to infer rough lower limits on X_{F} in the outer discs. The radial variation of X_{F} with respect to the mean value in the inner disc, X_{FM} , is a simple variation on expression (C1):

$$\frac{X_{\text{F}}}{X_{\text{FM}}} = \frac{N_{\text{d}}(\text{H}) - N(\text{H I})}{2 N_{\text{x}}(\text{H}_2)}, \quad (\text{C2})$$

where $N_{\text{x}}(\text{H}_2)$ is the molecular gas column density estimated from a spatially constant X-factor: $N_{\text{x}}(\text{H}_2) \equiv X_{\text{FM}} I(\text{CO})$. This measure of the column density is very noisy in these observations of the outer discs and a useful proxy is necessary for inferring some crude lower limit on $X_{\text{F}}/X_{\text{FM}}$ for radii beyond 8 kpc. The proxy adopted for $2N_{\text{x}}(\text{H}_2)$ was a pessimistic estimate of the 3σ uncertainties of the gas column density, i.e. $3[\sigma(N(\text{H I})) + 2\sigma(N_{\text{x}}(\text{H}_2))]$. This gives the most conservative (i.e. least extreme) estimates for $X_{\text{F}}/X_{\text{FM}}$ beyond a radius of 8 kpc. For M 51, even in the high- κ_{ν} (low- κ_{ν}) case, the above proxy yields lower limits to X_{F} of roughly 1 to $30 X_{20}$ ($7\text{--}10^3 X_{20}$) depending on the radius in the outer disc. For M 83, even in the high- κ_{ν} (low- κ_{ν}) case, these rough lower limits to X_{F} are $0.3\text{--}20 X_{20}$ ($2\text{--}600 X_{20}$).

APPENDIX D: CONSIDERATIONS OF THE COMPARISON BETWEEN SFR AND GAS SURFACE DENSITIES

The logarithm of the SFR surface density is plotted against that of the gas surface density for different tracers of that gas surface density. The plots of the surface densities of SFR versus gas are displayed in Fig. 16. The uncertainty of each fitted slope is the formal error of the fit and was scaled by the square root of the reduced chi square in order to give a more conservative and more realistic estimate of this formal error. One important consideration in these fits is that the spatial resolution of the M 51 data is limited to 38 arcsec, because of the *Spitzer* 160 μm resolution. For M 83, this resolution is 55 arcsec resolution due to the CO observations. Given that each point in these plots represents a single pixel, the effective number of independent points in each plot is the number of pixels divided by 20 for M 51 and by 55 for M 83.

Nevertheless, the linear resolution on each galaxy is about the same, with 1.5 kpc for M51 and 1.2 kpc for M83.

APPENDIX E: CONSIDERATIONS FOR CREATING SFR SURFACE DENSITY MAPS

The SFR surface density maps were created from images in the H α line and in the 24 μm continuum. According to Calzetti et al. (2007), the SFR, \dot{M}_* , within a source is given by

$$\dot{M}_* = a[L(\text{H}\alpha) + b \nu L_\nu(24 \mu\text{m})]. \quad (\text{E1})$$

$L(\text{H}\alpha)$ is the integrated luminosity of the H α line in erg s^{-1} ; ν is the frequency corresponding to the wavelength of 24 μm ; $L_\nu(24 \mu\text{m})$ is the luminosity of the 24 μm continuum per unit frequency bandwidth in $\text{erg s}^{-1} \text{Hz}^{-1}$; $a = 5.3 \times 10^{-42} M_\odot \text{yr}^{-1} (\text{erg s}^{-1})^{-1}$; $b = 0.031$; \dot{M}_* is in units of $M_\odot \text{yr}^{-1}$. If Σ_{SFR} is the SFR surface density, then $\Sigma_{\text{SFR}} = \dot{M}_*/(\Omega_s D^2)$, where Ω_s is the source solid angle and D is the source distance. Using $L = 4\pi D^2 F$, where F is the source flux, and also using $I = F/\Omega_s$, with I as the source surface brightness, results in

$$\Sigma_{\text{SFR}} = 4\pi a[I(\text{H}\alpha) + b c_\nu I_\nu(24 \mu\text{m})]. \quad (\text{E2})$$

For Σ_{SFR} in units of $M_\odot \text{yr}^{-1} \text{kpc}^{-2}$, $I(\text{H}\alpha)$ in units of $\text{erg s}^{-1} \text{cm}^{-2} \text{sr}^{-1}$, and $I_\nu(24 \mu\text{m})$ in units of MJy sr^{-1} , we have $a = 5.1 \times 10^1 M_\odot \text{yr}^{-1} \text{kpc}^{-2} (\text{erg s}^{-1} \text{cm}^{-2} \text{sr}^{-1})^{-1}$ and $c_\nu = 1.249 \times 10^{-4} \text{erg s}^{-1} \text{cm}^{-2} \text{sr}^{-1} (\text{MJy sr}^{-1})^{-1}$. The numerical value of b remains unchanged.

Specifically for M51 and M83, the SFR surface density maps were created by applying expression (E2) to the H α (i.e. the line-integrated images of Kennicutt et al. 2003; Blasco-Herrera et al. 2010) and the *Spitzer* 24 μm images (Dale et al. 2009) of both galaxies. The images were convolved to the appropriate resolution and rebinned to 9 arcsec pixels. The $I(\text{H}\alpha)$ image was converted to units of $\mu \text{erg s}^{-1} \text{cm}^{-2} \text{sr}^{-1}$. The final spatial resolution of the SFR surface density map for M51 is 38 arcsec, which is that of the *Spitzer* 160 μm image. For M83, this is 55 arcsec, which is that of the CO J = 1 \rightarrow 0 map. The negative tails of the histograms of pixel values provided estimates of the noise levels.

APPENDIX F: RELATIONSHIP BETWEEN THE X-FACTOR AND THE SFR SURFACE DENSITY

The physical models of Clark & Glover (2015) explore the possible dependence of the X-factor on the SFR. In their section 5.1, their results can be parametrized in the form

$$X_F = k_{\text{xs}} \Sigma_{\text{SFR}}^\gamma, \quad (\text{F1})$$

where they find $\gamma \simeq 0.5$. With the results of the fits illustrated in Fig. 16, we can test this result as well as specify the parameter k_{xs} .

We repeat the derivation in section 5.1 of Clark & Glover (2015) in more detail and adopt their nomenclature where applicable. If CO (1–0) is used to estimate the H₂ surface density, Σ_{mol} , then we let $\Sigma_{\text{mol}}(\text{CO})$ represent this surface density as estimated by CO. Then

$$\Sigma_{\text{SFR}} = k_{\text{sco}} [\Sigma_{\text{mol}}(\text{CO})]^{N_{\text{obs}}} = k_{\text{sco}} [X_{\text{FM}} I(\text{CO})]^{N_{\text{obs}}}, \quad (\text{F2})$$

where N_{obs} is the power-law index that is observed when using CO-derived molecular gas surface densities and X_{FM} is the mean X-factor value chosen for the galaxy observed.⁴ Using a tracer that supposedly gives the ‘true’ Σ_{mol} , would have a similar relation:

$$\Sigma_{\text{SFR}} = k_{\text{sh}} [\Sigma_{\text{mol}}]^{N_{\text{act}}} = k_{\text{sco}} [X_{\text{F}} I(\text{CO})]^{N_{\text{act}}}, \quad (\text{F3})$$

where N_{act} is the actual correct power-law index and X_{F} is the correct point-by-point value of the X-factor. Combining expressions (F2) and (F3) so as to eliminate $I(\text{CO})$ and comparing with (F1) yields

$$k_{\text{xs}} = X_{\text{FM}} \frac{k_{\text{sco}}^{1/N_{\text{obs}}}}{k_{\text{sh}}^{1/N_{\text{act}}}}, \quad (\text{F4})$$

and also,

$$\gamma = \frac{1}{N_{\text{act}}} - \frac{1}{N_{\text{obs}}}. \quad (\text{F5})$$

In the context of the current work, we assume that the dust tracers of the molecular gas give N_{mol} (i.e. $N_{\text{d}}(\text{H}) - N(\text{H}\text{I})$). So the fits in the second panels of Fig. 16 yield the parameters k_{sh} and N_{act} for M51 and M83. The fits in the fourth panels provide k_{sco} and N_{obs} . The uncertainties of these quantities are likely dominated by systematics. This means that k_{xs} has a roughly 50 per cent uncertainty. The effects of systematics on γ are difficult to estimate. The formal uncertainties of γ using the current data are 3–4 per cent.

With the above in mind, expressions (F4) and (F5) applied to the current data yield $k_{\text{xs}} = 0.18$ and $\gamma = -0.38$ for M51 and 0.79 and -0.25 for M83. Note that these values for k_{xs} give X_{F} in X_{20} units.

⁴ Note that the value chosen for X_{FM} is unimportant because it is the $k_{\text{sco}} X_{\text{FM}}^{N_{\text{obs}}}$ combination that matters.

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